

THE *SPITZER* c2d SURVEY OF LARGE, NEARBY, INTERSTELLAR CLOUDS. VII. OPHIUCHUS OBSERVED WITH MIPS

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Received 2007 January 10; accepted 2007 September 18

ABSTRACT

We present maps of 14.4 deg² of the Ophiuchus dark clouds observed by the *Spitzer Space Telescope* Multiband Imaging Photometer for *Spitzer* (MIPS). These high-quality maps depict both numerous point sources and extended dust emission within the star-forming and non–star-forming portions of these clouds. Using PSF-fitting photometry, we detect 5779 sources at 24 μ m and 81 sources at 70 μ m at the 10 σ level of significance. Three hundred twenty-three candidate young stellar objects (YSOs) were identified according to their positions on the MIPS/2MASS *K* versus *K* – [24] color-magnitude diagrams, as compared to 24 μ m detections in the SWIRE extragalactic survey. We find that more than half of the YSO candidates, and almost all those with protostellar Class I spectral energy distributions, are confined to the known cluster and aggregates.

Subject headings: infrared: stars — ISM: clouds — stars: formation

1. INTRODUCTION

The Ophiuchus molecular cloud complex is one of the most well-known regions of nearby active star formation. For the Lynds 1688 cloud alone, SIMBAD lists 770 references as of summer 2006. The clouds ($\alpha = 16^{\text{h}}28^{\text{m}}06^{\text{s}}$, $\delta = -24^{\circ}32'30''$ [J2000.0]) are considered to be at a distance of 125 ± 25 pc (de Geus et al. 1989), although more distant estimates are seen in the literature (145 pc, de Zeeuw et al. 1999; 160 pc, Chini 1981). The stellar population of the Ophiuchus clouds has been the subject of many imaging surveys (Grasdalen et al. 1973; Vrba et al. 1975; Wilking & Lada 1983; Lada & Wilking 1984; Rieke et al. 1989; Strom et al. 1995; Barsony et al. 1989, 1997, 2005; Greene & Young 1992; Comeron et al. 1993; Bontemps et al. 1998; Allen

et al. 2002; Duchêne et al. 2004; etc.) and spectroscopic surveys (Bouvier & Appenzeller 1992; Greene & Meyer 1995; Martin et al. 1998; Wilking et al. 1999, 2005; Luhman & Rieke 1999; Cushing et al. 2000; Natta et al. 2006; etc.). Many thousands of sources have been cataloged at near-IR wavelengths; however, the majority of these sources are extincted background stars (Barsony et al. 1997).

L1688 contains the most famous star-forming cluster in this region, with of order 100 stars < 1 Myr old, including several extremely young Class 0–I sources (Bontemps et al. 2001; Duchêne et al. 2004; Haisch et al. 2002; and many older studies). The vicinity of this cluster is covered in bright and dark nebulosity, and Herbig–Haro flows criss-cross the entire region (Gomez et al. 1998; Phelps & Barsony 2004). Most of the known young stellar objects (YSOs) are found within an embedded cluster only a few parsecs in extent, making Ophiuchus one of the higher density environments among the nearby low-mass star formation regions (Grasdalen et al. 1973; Wilking & Lada 1983, 1989; etc.). It is also a region forming intermediate-mass stars, bridging the gap between Taurus (a region of low-mass, isolated star formation) and Orion (a dense cluster with massive O stars). In Ophiuchus, the principal region of low-mass star formation is in the vicinity of several intermediate- to high-mass stars of the Sco OB1 association (de Zeeuw et al. 1999), and a number of A and B stars appear to be embedded within the cloud (Klose 1986; Devito & Hayward 1998).

The molecular material in the Lynds 1688 and 1689 clouds extends far beyond the boundaries of the extant surveys for YSOs. To the southeast of L1688 is the L1689 cloud, which extends to the east in the L1712 and L1729 streamers. Only a few isolated YSOs are known in this fairly massive cloud, leading Nutter et al. (2006) to dub it “the dog that didn’t bark.” Their submillimeter continuum mapping survey indicates that most of the mass of L1689 resides in distributed cloud material rather than dense cores, unlike L1688, and suggests that the proximity of L1688 to σ Sco explains its much greater star formation efficiency. Filamentary clouds (L1709, L1740, and L1744+L1755+L1765 = “Oph North”) also extend to the northeast of L1688. To date,

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TABLE 1
SUMMARY OF OBSERVATIONS (PROGRAM 177)

FIELD	MAP CENTER		FIRST-EPOCH AORKEY	SECOND-EPOCH AORKEY
	α (J2000.0)	δ (J2000.0)		
Ophi-mips1.....	16 45 50.0	-21 21 16.0	5745408	5747200
Ophi-mips2.....	16 43 24.0	-21 57 38.0	5745664	5747456
1689-mips1.....	16 43 16.0	-24 09 09.4	5748992	5753344
1689-mips2.....	16 38 27.0	-24 18 02.0	5749248	5753600
1689-mips3.....	16 35 12.0	-24 34 00.0	5749504	5753856
1688-mips1.....	16 33 33.0	-24 20 42.0	5757952	5766144
1688-mips2.....	16 33 01.5	-23 18 20.0	5758208	5766400
1688-mips3.....	16 31 59.2	-24 57 53.0	5758464	5766656
1688-mips4.....	16 31 00.0	-23 37 50.1	5758720	5766912
1688-mips5.....	16 30 03.0	-25 08 22.0	6605568	6605824
1688-mips6.....	16 27 56.5	-24 11 12.0	5759488	5767680
1688-mips7.....	16 25 46.0	-24 12 29.0	5759744	5676936
1688-mips8.....	16 24 09.0	-23 50 03.0	5760000	5768192
1688-mips9.....	16 21 52.5	-23 43 45.0	5760256	5768448
Oph-oc7.....	16 23 52.0	-21 57 38.0	5778432	5778688
Oph-oc8.....	16 23 12.0	-26 28 49.0	5778944	5779200

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds.

these clouds are lacking in high-resolution mid-IR surveys to identify and characterize their YSO content. Only the *IRAS* survey and the Two Micron All Sky Survey (2MASS) exist for these regions. The ISOCAM survey of Oph covered only 0.7 deg^2 (Bontemps et al. 1998). Those authors claim that most of the YSOs outside of this central region are weak-line and post-T Tauri stars, and indeed, this is generally true based on only the *IRAS* and *ROSAT* studies of the region. The current study using the *Spitzer Space Telescope* is capable of detecting low-luminosity Class 0–III infrared sources throughout a much larger surveyed area.

The Ophiuchus molecular cloud complex is one of five star-forming regions observed with *Spitzer* as a part of the “Molecular Cores to Planet-forming Disks” (c2d) Legacy project (Evans et al. 2003). The goal of this survey is to characterize the circumstellar material of young stars and substellar objects in these clouds. This paper is one in a series which present the c2d large-cloud mapping results made with *Spitzer*’s Infrared Camera (IRAC) and Multiband Imaging Photometer for *Spitzer* (MIPS). Others in the series report results for Chamaeleon II (MIPS, Young et al. 2005; IRAC, Porras et al. 2007), Lupus (MIPS, Chapman et al. 2007), Perseus (MIPS, Rebull et al. 2007; IRAC, Jørgensen et al. 2006), and Serpens (IRAC, Harvey et al. 2006; MIPS, Harvey et al. 2007). These papers endeavor to present the results for each cloud and instrument in a standard format. Although the current work principally discusses the Ophiuchus MIPS maps, data from IRAC are also integrated into sections regarding source classification and individual objects. The IRAC results for Ophiuchus will be presented in L. E. Allen et al. 2008, in preparation.

The c2d MIPS survey of Ophiuchus attempts to determine the importance of a more distributed mode of star formation in the less-studied areas outside the Oph core by covering more than 13.8 deg^2 in the region where $A_V \geq 5$ at sensitivity levels considerably better than that of ISOCAM at mid-IR wavelengths and many orders of magnitude better than *IRAS* and ISOPHOT at far-IR wavelengths. The three MIPS filters at 24, 70, and $160 \mu\text{m}$ are ideally suited for linking the circumstellar environment of individual stars to the interstellar environment of the cloud. The MIPS $24 \mu\text{m}$ band is especially useful for identifying stars with infrared

excess, since stellar photospheric emission has decreased by 2 orders of magnitude from the blackbody peak, increasing the contrast with peak emission from cool dust. Thus, stars with warm young disks are generally extremely bright at $24 \mu\text{m}$ relative to stars without disk emission. MIPS $70 \mu\text{m}$ emission is also seen in the brighter envelopes and disks and in interstellar regions where dust is heated by young stars. MIPS $160 \mu\text{m}$ traces cool dust in both circumstellar and interstellar regions, similar to the *IRAS* $100 \mu\text{m}$ band. The combination of MIPS data with surveys at other wavelengths enables detailed studies of YSO spectral energy distributions (SEDs) and the cloud environment. The most recent millimeter continuum maps of this region have been presented by Young et al. (2006), who used the Caltech Submillimeter Observatory (CSO) Bolocam to map the region observed by *Spitzer*. High-quality millimeter molecular-line maps of Ophiuchus also covering the *Spitzer* survey are available from the COMPLETE team (Ridge et al. 2006).

In this paper we first present the *Spitzer* MIPS observations of Ophiuchus (§ 2). Following the observations section, we discuss results, including multiwavelength images (§ 3.1), MIPS source counts (§ 3.2), band-merging with IRAC and 2MASS (§ 3.3), and a variety of color-color and color-magnitude diagrams (§ 3.4). Then we investigate recovery of *IRAS* objects (§ 3.5) and variability at $24 \mu\text{m}$ on timescales of 3–8 hr in § 3.6. We highlight the SEDs and imaging of several individual objects in § 4.1. We then discuss the distribution of infrared excess sources in Ophiuchus and comment on the degree of clustering found throughout the region (§ 4.2). Finally, we summarize our main points in § 5.

2. OBSERVATIONS AND DATA REDUCTION

2.1. MIPS

The MIPS observations of Ophiuchus were conducted on 2004 March 18–19 and covered 13.8 deg^2 . Two off-cloud fields totalling 0.6 deg^2 were observed with the intention of characterizing the background star counts. The center of this large map is roughly at $(\alpha, \delta)(\text{J2000.0}) = (16^{\text{h}}33^{\text{m}}18.0^{\text{s}}, -24^{\circ}15')$, or $(l, b) = (354^{\circ}, +16.0^{\circ})$, or ecliptic coordinates $(251^{\circ}, -2.5^{\circ})(\text{J2000.0})$. The area mapped with MIPS completely covers the c2d IRAC

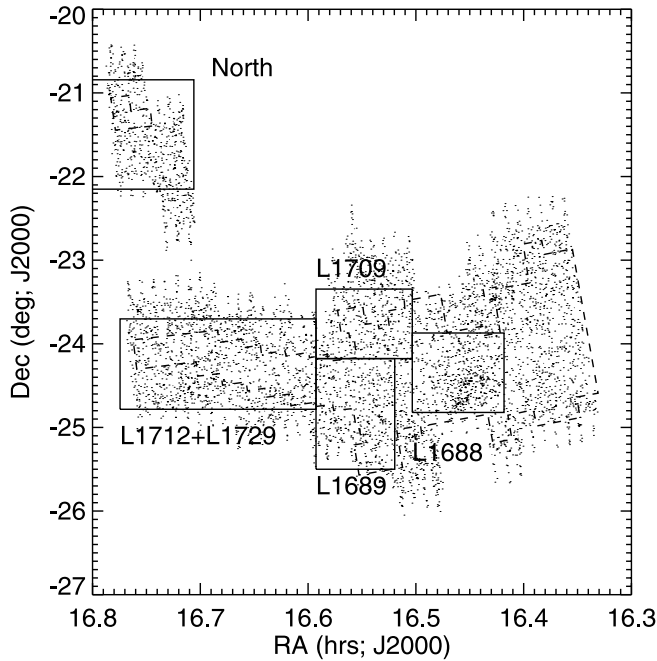


FIG. 1.—Location of MIPS coverage (points are MIPS-24 detections), with the region of IRAC coverage (dashed line) indicated. The smaller squares (solid line) indicates the regions defined to be L1688, L1689, and Oph North.

map regions in Ophiuchus, which included all the cloud area above the $A_V = 5$ contour. The combination of the cloud being extended perpendicular to the MIPS scan direction, and the fixed scan length options available for use with MIPS, dictated that the area covered by the Ophiuchus MIPS maps be more than twice the area mapped with IRAC.

These observations were part of *Spitzer* program 177; the AORKEYs appear in Table 1. Fast-scan maps were obtained at two separate epochs separated by a few hours. At each epoch, the spacing between adjacent scan legs was $240''$. The second-epoch observation was offset $125''$ in the cross-scan direction from the first to fill in the $70 \mu\text{m}$ sky coverage, which was incomplete at each individual epoch. Furthermore, the second-epoch scan was also offset $80''$ from the first in the scan direction to maximize $160 \mu\text{m}$ map coverage. These mapping parameters resulted in every part of the map being imaged at two epochs at $24 \mu\text{m}$ and only one epoch at 70 and $160 \mu\text{m}$, with total integration times of 30, 15, and 3 s at each point in the map (respectively). Even with this strategy, the $160 \mu\text{m}$ maps have some gaps, which makes point-source photometry especially uncertain at this wavelength. In addition, the $160 \mu\text{m}$ array is saturated in the region at the center of the L1688 cloud.

Figure 1 shows the regions of three-band coverage with MIPS and of four-band coverage with IRAC. In addition to the c2d data, smaller MIPS maps of Ophiuchus were obtained by *Spitzer* Guaranteed Time Observers. To maintain uniform sensitivity in our maps, these data were not incorporated into the current study.

The observations were conducted at two separate epochs separated by about 3–8 hr to permit asteroid removal in this low ecliptic latitude ($\sim 0^\circ$) field. Study of the asteroid population in Taurus at a similar ecliptic latitude (0° – 7°) indicates there are roughly 250 asteroids deg^{-2} with $24 \mu\text{m}$ flux densities $\geq 1.6 \text{ mJy}$ (Stapelfeldt et al. 2006). By extracting sources separately from the two epochs and comparing the source lists, transient sources were eliminated from our MIPS catalog. This observation technique also enables a study of $24 \mu\text{m}$ source variability on a time-scale of several hours (see § 3.6).

TABLE 2
BOX REGIONS USED TO DEFINE INDIVIDUAL CLOUDS

CLUSTER	SOUTHEAST BOX CORNER		NORTHWEST BOX CORNER	
	α (J2000.0)	δ (J2000.0)	α (J2000.0)	δ (J2000.0)
L1688.....	16 30 12.0	–24 49 15.6	16 25 05.5	–23 49 15.6
L1689.....	16 35 33.6	–25 30 00.0	16 31 10.6	–24 10 37.2
L1709.....	16 35 33.6	–24 10 37.2	16 30 12.0	–23 20 45.6
L1712 + L1729.....	16 46 28.1	–24 46 58.8	16 35 33.6	–23 42 07.2
North.....	16 48 10.3	–22 09 03.6	16 42 21.6	–20 50 34.8

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds.

The MIPS observations of Ophiuchus included two small off-cloud regions with total area of 0.66 deg^2 , in addition to the 13.8 deg^2 of the main map. These were intended to provide statistics on the stellar and extragalactic backgrounds. However, these regions were found to be too small and sparse to yield reliable statistics. In addition, one of the fields inadvertently included the large globular cluster M4, which severely biases its source counts. Thus, we did not utilize these off-cloud regions; instead, we compare our statistics to the 6.15 deg^2 ELAIS N1 field mapped by the *Spitzer* Wide-Area Infrared Extragalactic (“SWIRE”) survey (Shupe et al. 2008), which was rereduced and analyzed using the c2d pipeline. In the low ecliptic latitude Ophiuchus region, MIPS source counts are dominated by asteroids and the extragalactic background rather than stellar photospheres, and thus, a comparison to the SWIRE field is invaluable despite its considerably different Galactic latitude. However, we do include the off-cloud fields (OC1 and OC2) in the table of statistics (see § 3.2).

We started with the *Spitzer* Science Center (SSC)-pipeline-produced basic calibrated data (BCDs), version S11.4. Each MIPS channel was then processed differently, and therefore, each is discussed separately below. Additional detail on the data reduction process can be found in Young et al. (2005), Chapman et al. (2007), and Rebull et al. (2007).

Standard c2d pipeline processing was used for MIPS-24 (Young et al. 2005; Evans 2003). After some corrections for nonuniform readout variations (jailbars) were applied to individual frames (BCDs), these data were transformed into single-epoch mosaics using the SSC MOPEX software (following Chapman et al. 2007). Source lists were extracted using c2dphot (Harvey et al. 2006) to measure flux densities using point-spread function fitting. The final working 2005 catalog as assembled consists of all detections at MIPS-24 with high-quality c2d catalog detection flags (Evans et al. 2007, the c2d Data Delivery document¹⁵) and detections at both epochs. While resulting in a shallower survey than would be possible using other combinations of flags, this ensures that no asteroids are included in the final catalog. The zero point used to convert these flux densities to magnitudes was 7.14 Jy , based on extrapolation from the Vega spectrum as published in the MIPS Data Handbook. There were 5779 total point sources detected at $24 \mu\text{m}$ meeting our criteria, ranging from 0.672 to 3910 mJy . Note that the spatial resolution of MIPS-24 is $\sim 6''$, with a $2.55''$ pixel size.

For the MIPS-70 data we started with the automated pipeline-produced BCDs, both the filtered and unfiltered products. Then individual BCDs were mosaicked into one filtered and one unfiltered mosaic using the SSC software MOPEX (Makovoz

¹⁵ See http://data.spitzer.caltech.edu/popular/c2d/20071101_enhanced_v1/Documents/c2d_del_document.pdf.

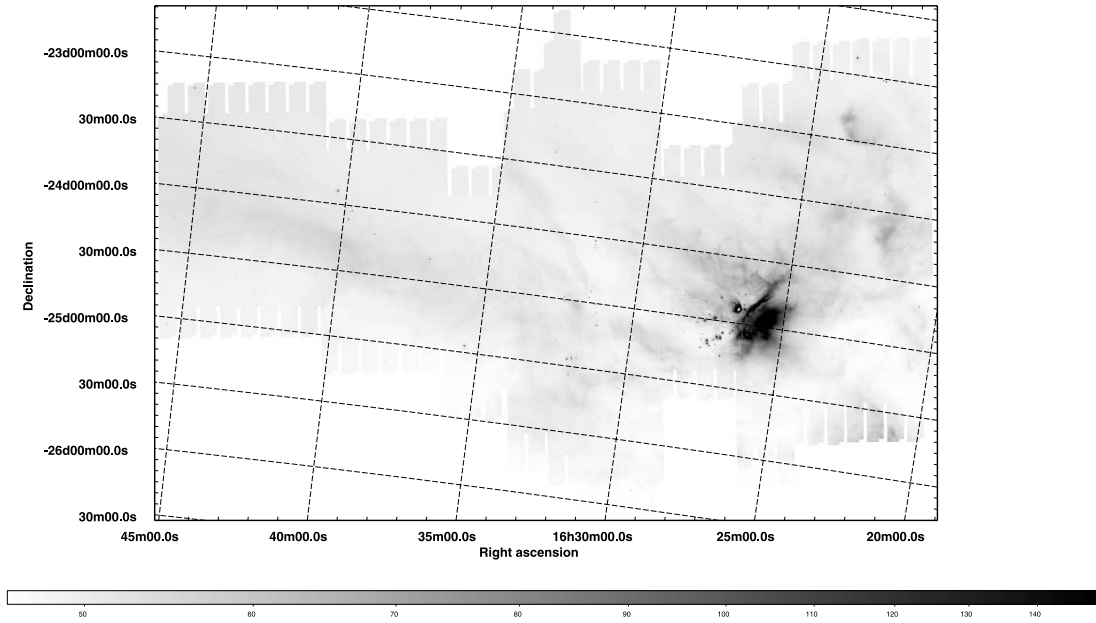


FIG. 2.—*Spitzer* MIPS 24 μm mosaic of Ophiuchus clouds.

et al. 2006). We performed source detection and extraction using MOPEX/APEX on the entire mosaic. The source list was cleaned for instrumental artifacts via manual inspection of the 70 μm image and comparison to the 24 μm image; e.g., if there was some question as to whether a faint object seen at 70 μm was real or an instrumental artifact, and comparison to the 24 μm image revealed a 24 μm source, then the 70 μm object was judged to be a real source. About 80% of the 70 μm objects had identifiable counterparts at 24 μm . The zero point used to convert the flux densities to magnitudes was 0.775 Jy, based on extrapolation from the Vega spectrum as published in the MIPS Data

Handbook. There were 81 total point sources detected at 70 μm . There were 19 objects brighter than 4 Jy, for which we employed aperture photometry with a 32'' aperture and a 17% aperture correction. Note that the spatial resolution of MIPS-70 is $\sim 18''$.

The 160 μm MIPS band presented challenges for point-source extraction due to scattered gaps in the map coverage. A small amount of smoothing was applied to the data in order to fill gaps. Then APEX was used to find sources throughout the image. Because of the large saturated region in L1688 and the highly structured nature of the bright diffuse emission, few isolated point sources were identified in our study. The zero point used to

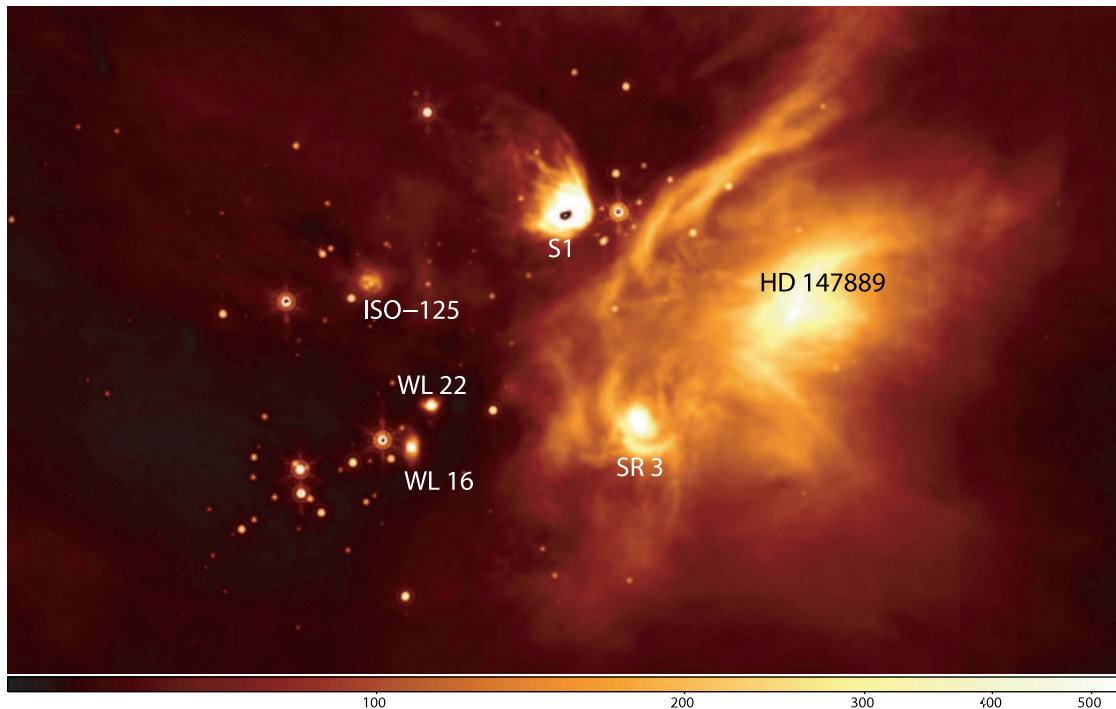


FIG. 3.—MIPS 24 μm image: central $1^\circ \times 0.5^\circ$ region of Lynds 1688.

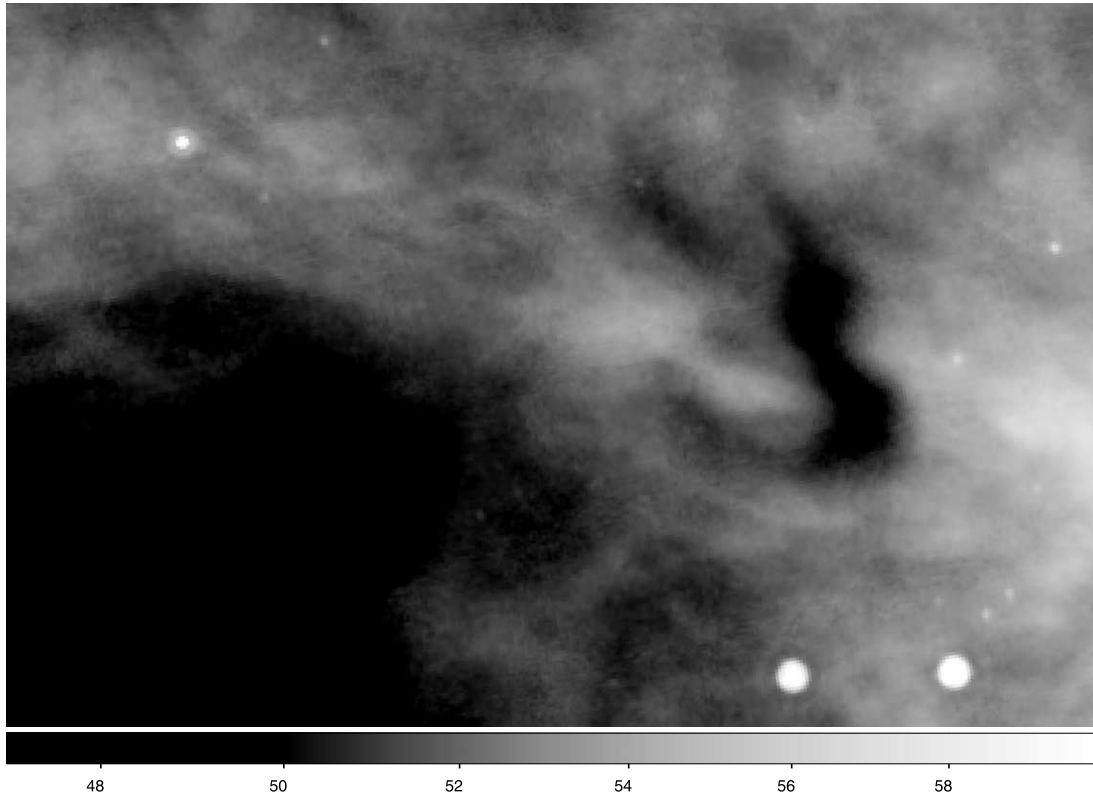


FIG. 4.—Dark clouds at $24\ \mu\text{m}$ found near $16^{\text{h}}28^{\text{m}}31^{\text{s}}$, $-24^{\circ}18'40''$. The image spans $15' \times 10'$, with north up and east to the left.

convert the flux densities to magnitudes was $0.159\ \text{Jy}$, based on extrapolation from the Vega spectrum as published in the MIPS Data Handbook. There were eight total point sources detected at $160\ \mu\text{m}$, ranging from $876\ \text{mJy}$ to $19\ \text{Jy}$. About 60% of the $160\ \mu\text{m}$ objects had identifiable counterparts at 24 and $70\ \mu\text{m}$. The extreme brightness and saturation of the $160\ \mu\text{m}$ diffuse emission made detection of $160\ \mu\text{m}$ sources impossible in the most densely populated regions of the cluster. Finally, note that the spatial resolution of MIPS-160 is $\sim 36''$.

Although this paper is primarily about point sources, there is extensive extended emission in all three MIPS channels throughout the MIPS maps. The diffuse emission ranges in brightness from 50 to $570\ \text{MJy sr}^{-1}$ at $24\ \mu\text{m}$, where the low value is set by the overall zodiacal emission in this low ecliptic latitude region. The $70\ \mu\text{m}$ extended emission ranges in surface brightness from ~ 30 to $5000\ \text{MJy sr}^{-1}$, while the $160\ \mu\text{m}$ extended emission varies from 45 to $400\ \text{MJy sr}^{-1}$ (with the central region of L1688 being saturated). This extreme background variation unfortunately

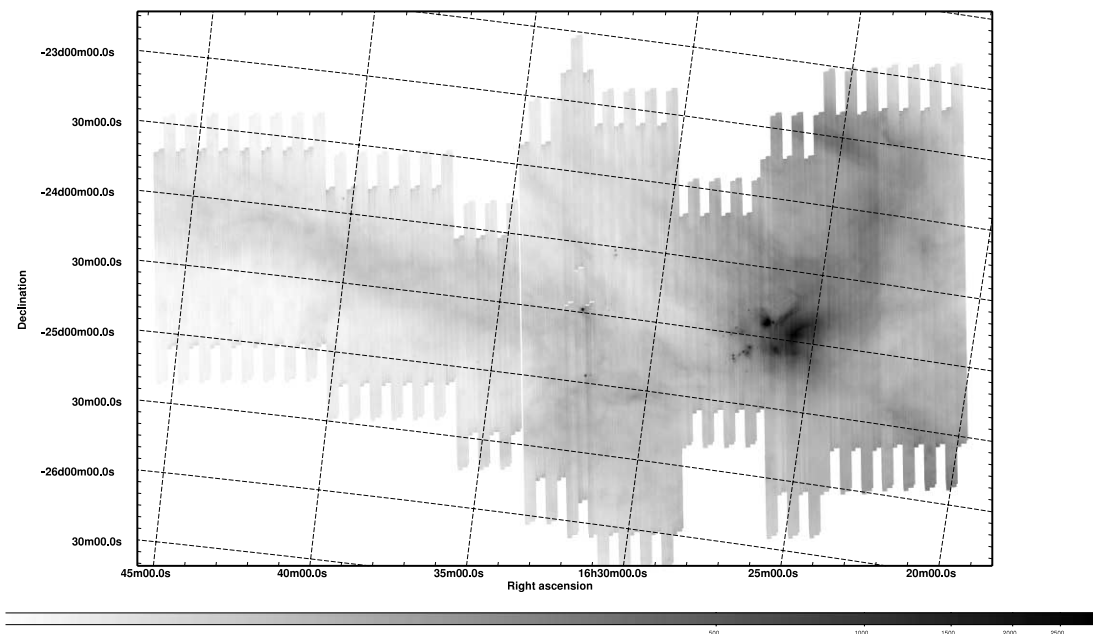


FIG. 5.—*Spitzer* MIPS $70\ \mu\text{m}$ mosaic of Ophiuchus clouds.

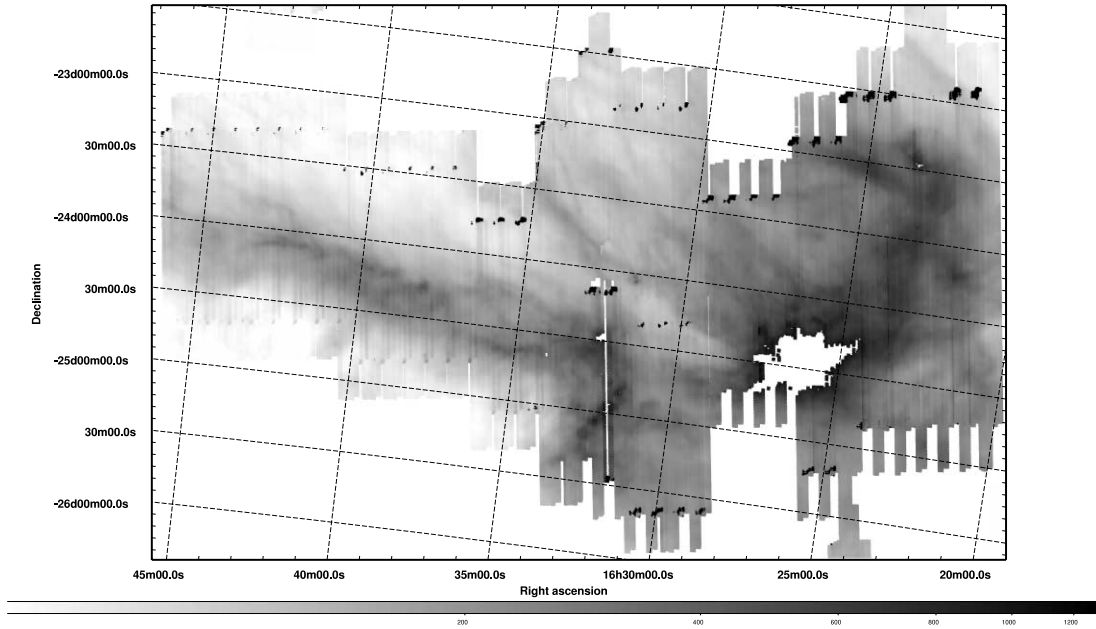


FIG. 6.—*Spitzer* MIPS $160\ \mu\text{m}$ mosaic of Ophiuchus clouds. The central region of the cloud is saturated in our maps.

results in highly nonuniform point source detection sensitivity in our survey. Our region of worst sensitivity coincides with the central core of L1688, where the most intense past studies have been performed. Thus, our *Spitzer* c2d survey should be considered complementary to past *Infrared Space Observatory (ISO)* and ground-based IR studies in Oph.

Because a large cloud complex is encompassed by our maps and the population of the individual clouds is known to have substantially different characteristics, we have chosen to separate out the properties of stars in L1688, L1689, L1709, and the

L1712+L1729 filaments (see Fig. 1). In addition, we have separated out the properties of the disconnected northern filaments (L1744+L1755+L1765), which we identify as Oph North. The above regions were defined as the rectangles shown in Figure 1, with the corners given in Table 2.

2.2. CSO *Sharc II*

Several L1709 sources were observed at $350\ \mu\text{m}$, using *Sharc II* at the CSO on 2007 April 23, 25, and 26. *Sharc II* is a 12×32 bolometer array, with a pixel size of $4.85''$ and a field of view of

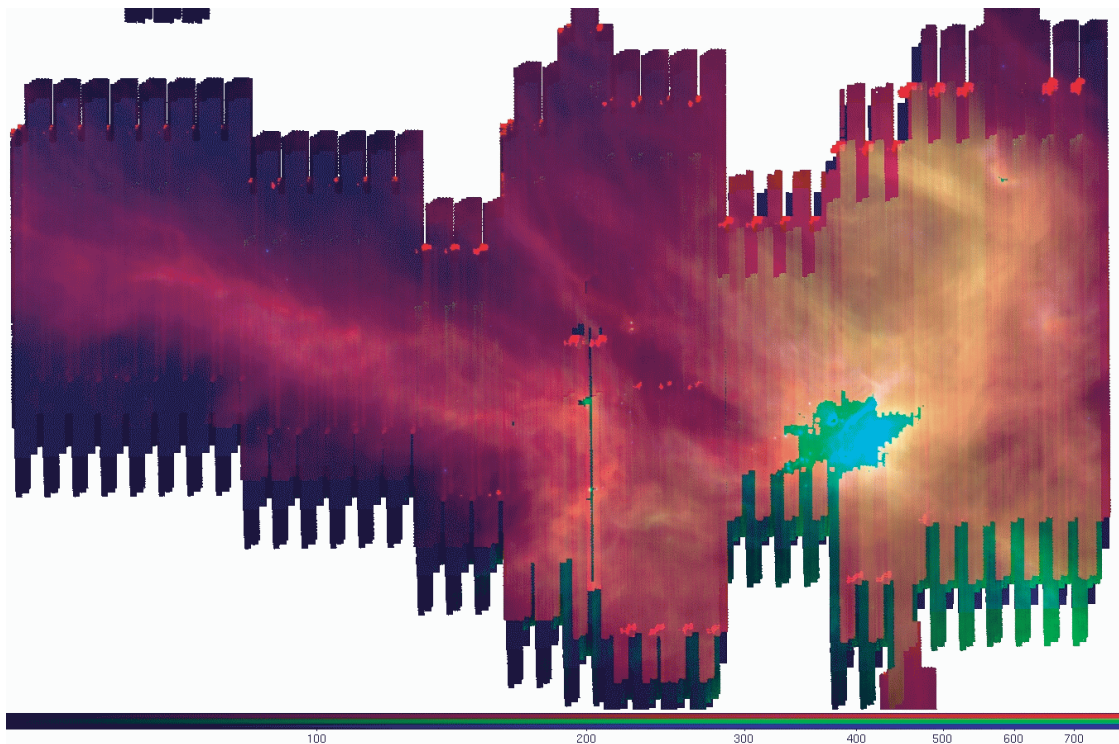


FIG. 7.—A 24 (blue), 70 (green), and $160\ \mu\text{m}$ (red) mosaic of Ophiuchus clouds.

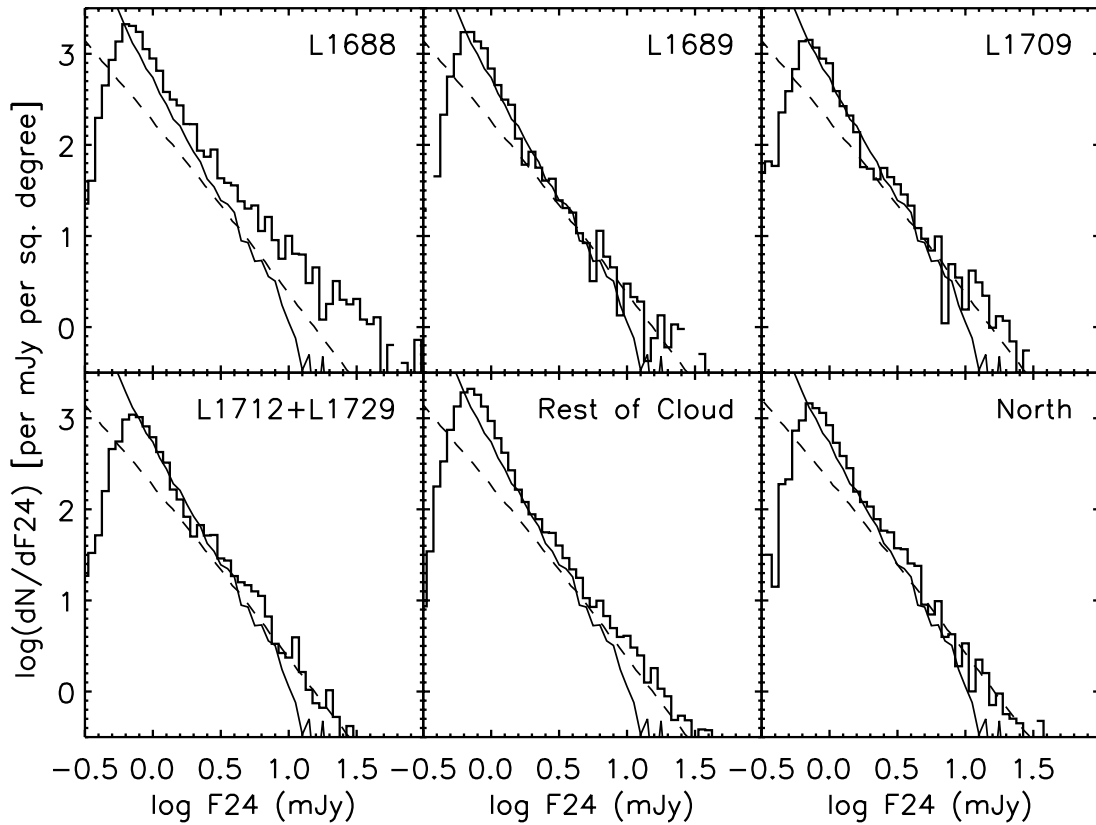


FIG. 8.—Ophiuchus differential number counts at $24 \mu\text{m}$. The solid line shows the SWIRE ELAIS N1 extragalactic number counts, and the dashed line shows the Wainscoat model prediction of Galactic star counts.

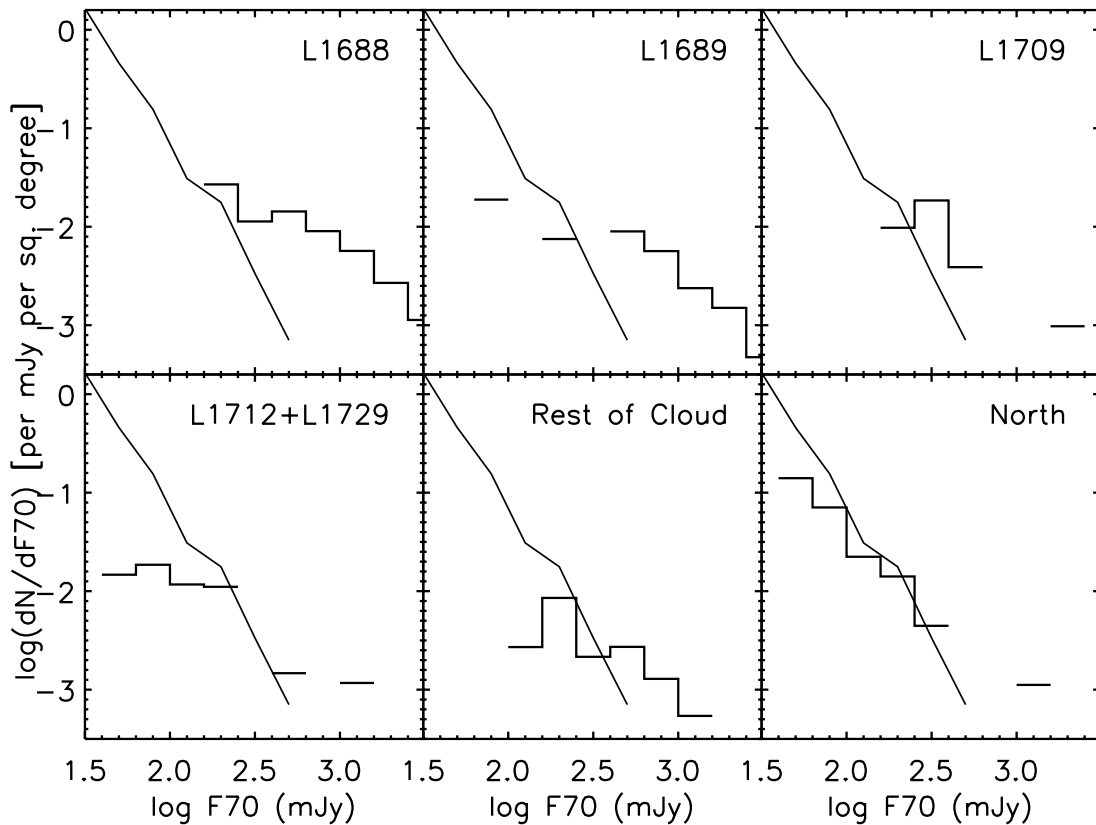


FIG. 9.—Ophiuchus differential number counts at $70 \mu\text{m}$. The solid line shows the SWIRE ELAIS N1 extragalactic number counts.

TABLE 3
DETECTED SOURCES

Item	Overall	IRAC Coverage	L1688	L1689	L1709	L1712+L1729	Rest of Cloud	North	Off-Cloud
No. 24 μm objects.....	5779	2608	569	490	377	1082	2555	706	355
No. 70 μm objects.....	81	31	24	12	5	12	14	13	2
No. 160 μm objects.....	8	6	1	1	2	1	1	2	0
No. 24 and 70 μm objects.....	50	31	13	6	4	5	11	11	0
No. 24, 70, and 160 μm objects.....	5	4	1	1	1	1	0	1	0
No. 24 μm and <i>K</i> objects.....	2590	1197	297	195	149	539	1102	308	177
No. 24 μm and any IRAC objects.....	2306	2306	425	351	144	379	895	112	0
No. 70 μm and any IRAC objects.....	42	30	23	7	4	3	4	1	0
No. 160 μm and any IRAC objects.....	6	6	1	1	2	1	0	1	0

$2.6' \times 0.97'$. The Dish Surface Optimization System was used throughout the run, providing a stable beam, with a FWHM of $\sim 8''$ at $350 \mu\text{m}$. Submillimeter conditions were excellent throughout the run, with τ_{225} ranging from 0.03 to 0.045. Target observations were interspersed with observations of IRAS 16293–2422 and Callisto, which were used for both pointing and flux correction. The data were reduced using the CSO data pipeline CRUSH and relative aperture photometry on the sources obtained. The measurement uncertainties include the flux calibration uncertainty, which is typically on the order of 20%.

3. RESULTS

3.1. Multiwavelength Images

The MIPS 24 μm mosaic (Fig. 2) shows a great deal of diffuse nebulosity and many bright point sources. The central region of L1688 (Fig. 3) is dominated by nebulosity near the early-type stars (S1 and SR 3) embedded in the cloud. The S1 source is saturated in the MIPS 24 μm image, but the peak appears extended, and it is surrounded by a complex cometary nebula opening toward the northeast. The B star SR 3 also appears as a resolved nebulosity at 24 μm , as do the more embedded early-type objects WL 16 and WL 22. The WL 22 nebula has a central core with a FWHM of $18''$ and an outer cometary morphology extending to radii of $30''$. The WL 16 nebula is more elongated, $2' \times 35''$ along a position angle (P.A.) of $10^\circ - 50^\circ$ from the major axis of the polycyclic aromatic hydrocarbon disk reported by Ressler & Barsony (2003). An elliptical ring of emission about $1'$ across is centered on a point source near the position of ISO–Oph 125 and ISO–Oph 124 (YLV 12B and 12A; spectral types F7 and M4; Greene & Meyer 1995; Luhman & Rieke 1999). Two other point sources within this ring are also detected at the positions of ISO–Oph 128 and ISO–Oph 129. South of these sources, the HD 147889 star appears only as a faint point source against some of the brightest parts of the central Oph nebulosity and shows a purely photospheric SED. This confirms that the large FIR excesses assigned to this source by the *IRAS* Point Source Catalog (PSC) are most likely due to interstellar emission, rather than circumstellar material. Ground-based photometry of this source (Wilking et al. 1989) also has suggested a lack of circumstellar material in this system. However, HD 148579 (B9 V), south of L1689, has a resolved 24 μm nebula, as does the nearby HD 148605 (B3 V). Interestingly, these stars, which are apparently adjacent to or embedded in cloud material, have not excited the interstellar nebulosity to the same degree associated with the early-type stars in the central cluster. The 24 μm mosaic also shows a surprising amount of “dark” nebulosity (i.e., absorbing perhaps 80% of the background illumination in the vicinity of the L1688 cluster and

extending toward L1689). One of the 24 μm dark cloudlets is shown in Figure 4.

The longer wavelength MIPS mosaics are shown in Figures 5 and 6. The 70 μm mosaic still shows striations from the detector history effects which are not entirely removed in the pipeline processing for each scan leg. Contributing to this appearance is the fact that the map was filled by interleaving observations at two different epochs. Despite the cosmetic imperfections, the MIPS 70 μm map clearly shows that the diffuse emission at this wavelength is brightest in the L1688 region and in the area extending from northwest to southwest of the central cluster. The COMPLETE CO maps of the Oph clouds (Ridge et al. 2006) show that this region of Oph West consists of generally tenuous cloud material with a few interspersed clumps. It is possible that these regions are bright in the long-wavelength MIPS bands because they are on a direct line of sight to the young B stars in the central cluster and are heated by them. S1, WL 16, WL 22, and YLV 12B continue to be nebulous at 70 μm as they are at 24 μm . There is also a small area of diffuse nebulosity adjacent to HD 148605. The sources IRAS 16293–2422 and IRAS 16288–2450 are also prominent at this wavelength. The 160 μm mosaic (Fig. 6) shows the same L1688 west dust emission features as the 70 μm image. However, the cold dust in L1689 and the L1712+1729 filaments becomes prominent at this wavelength. Young et al. (2006) showed that there are some weak millimeter sources in these filaments. The central region of L1688 is saturated in the MIPS 160 μm image. Figure 7 shows a three-color image with all three channels of MIPS included.

3.2. Source Counts

MIPS source counts in Ophiuchus are a combination of cloud members, foreground/background stars in the galaxy, and the extragalactic background. An excess of sources above comparison regions outside the clouds implies the presence of young stars with disks. Figure 8 shows the observed Ophiuchus 24 μm differential source counts in comparison to observed source counts from the SWIRE ELAIS N1 extragalactic field (*solid lines*) and to the prediction for Galactic star counts in the *IRAS* 25 μm band from the Wainscoat & Cowie (1992) model provided by J. Carpenter (2001, private communication; *dashed lines*). Aside from L1688, all of the other regions have negligible levels of excess 24 μm source counts, indicating that only small numbers of young stars are present in these fields. L1709 has a modest excess above 10 mJy. These plots indicate that only L1688 has a substantial young population relative to the various foreground/background sources expected in the field.

Figure 9 shows 70 μm source counts in Ophiuchus, along with extragalactic background counts from the SWIRE ELAIS N1

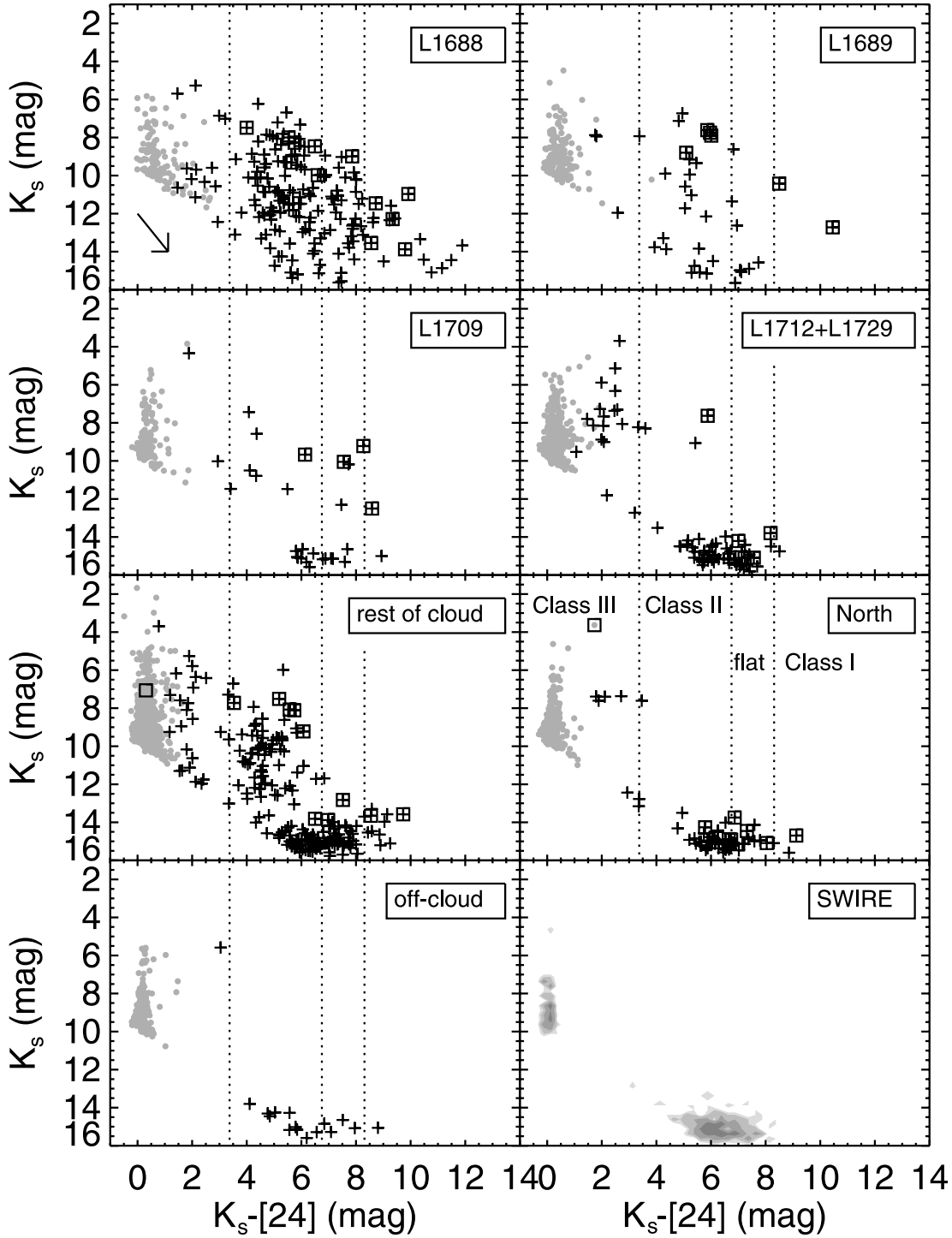


FIG. 10.— K_s vs. $K_s - [24]$ for Ophiuchus objects in the various subclusters as marked, the “rest of the cloud,” the off-cloud region, and SWIRE (contour plot, bottom right). Objects in SWIRE are expected to be mostly galaxies (objects with $K_s \geq 14$) or plain photospheres (objects with $K_s - [24] \leq 1$; marked as gray on the Oph region plots). Boxed sources are also detected at $70 \mu\text{m}$. The vertical dashed lines delineate (left to right) regions where Class I, flat-spectrum, Class II, and Class III infrared source models would reside.

field. At the low sensitivity of our $70 \mu\text{m}$ measurements, detection of normal stellar photospheres will be negligibly rare; thus, comparisons with Galactic star counts are not needed. Clearly, given the high backgrounds and small number of cataloged $70 \mu\text{m}$ sources, we are dealing with small-number statistics. However, the plots do show significantly more sources at bright flux density levels for the regions L1688, L1689, and possibly L1709 than would be found from the extragalactic background. For L1688, the

excess sources extend down to 200 mJy, while for the other two regions, the excess above background starts at about 300 mJy. In the extended cloud outside the clusters, a slight excess in source counts might be present above a flux density of 400 mJy. Note that the depth of $70 \mu\text{m}$ extractions is greatest in the Oph North field, where the diffuse background is lowest. In this field the $70 \mu\text{m}$ source counts follow the SWIRE survey curve to the point where the greater SWIRE sensitivity diverges from the c2d sensitivity.

TABLE 4
CLASSIFICATION BASED ON $K_s - [24]$

Item	Entire Cloud	L1688	L1689	L1709	L1712+L1729	Rest of Cloud	North	Off-Cloud
No. objects with K_s and $24 \mu\text{m}$	2590	297	195	149	539	1102	308	177
No. with $K_s - [24] > 2$, $K_s < 14$	323	161	27	13	18	96	8	2
No. with $K_s - [24] > 2$, $K_s < 14$, and Class I $K_s - [24]$ color	20 (6%)	12 (7%)	2 (7%)	1 (7%)	0 (0%)	5 (5%)	0	0
No. with $K_s - [24] > 2$, $K_s < 14$, and “flat” $K_s - [24]$ color	47 (14%)	33 (20%)	3 (11%)	4 (30%)	1 (5%)	5 (5%)	1 (12%)	0
No. with $K_s - [24] > 2$, $K_s < 14$, and Class II $K_s - [24]$ color	207 (64%)	100 (62%)	20 (74%)	7 (53%)	5 (27%)	73 (76%)	2 (25%)	1 (50%)
No. with $K_s - [24] > 2$, $K_s < 14$, and Class III $K_s - [24]$ color	49 (15%)	16 (9%)	2 (7%)	1 (7%)	12 (66%)	13 (13%)	5 (62%)	1 (50%)

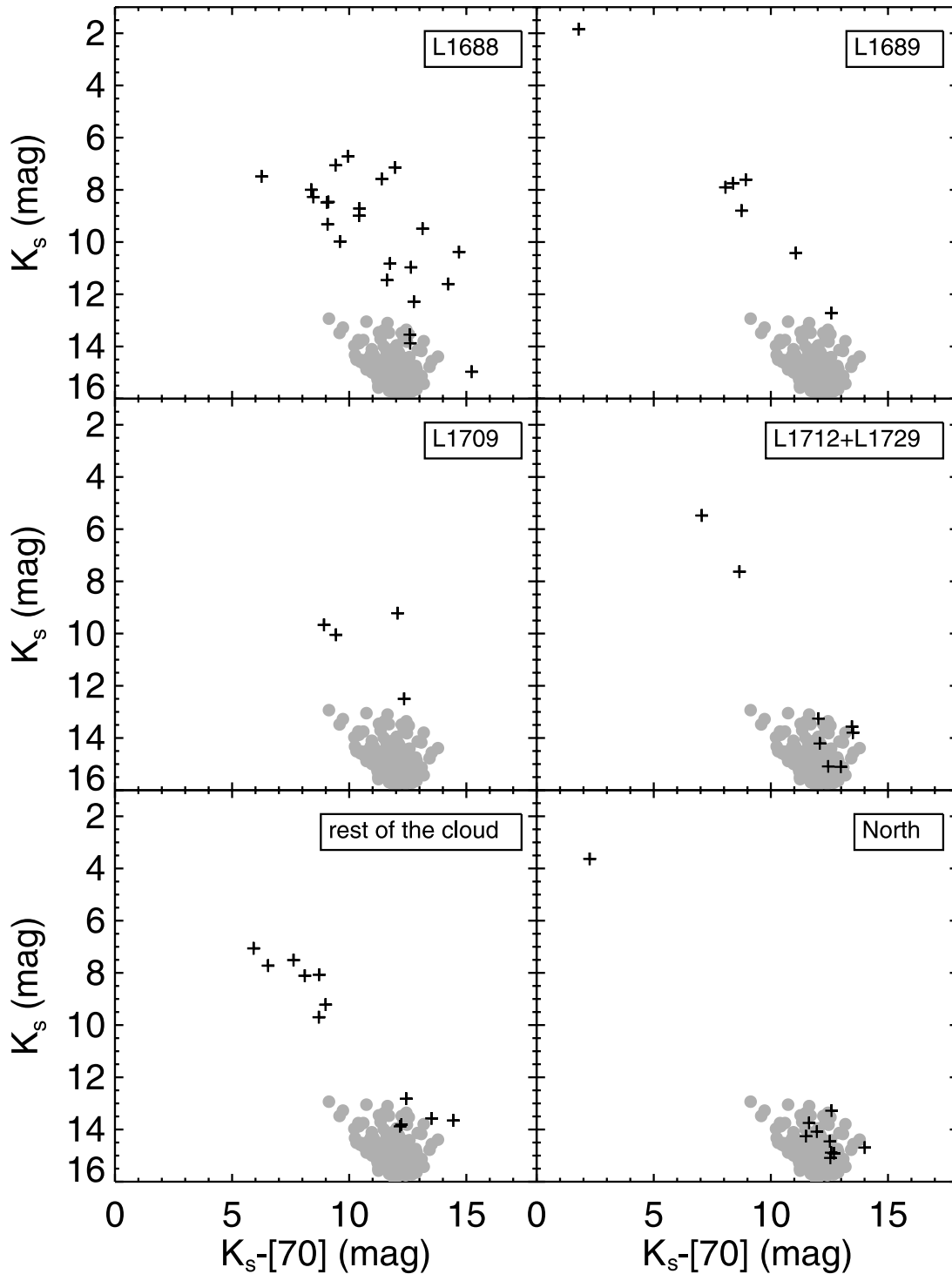


FIG. 11.— K_s vs. $K_s - [70]$ color-magnitude diagram for various regions in Oph (crosses), with data for SWIRE (gray circles) included for comparison. See text for details.

3.3. Band-Merging with IRAC and 2MASS

The IRAC data for the c2d Ophiuchus field (covering $\sim 6.8 \text{ deg}^2$) will be described in L. E. Allen et al. (2008, in preparation), and these data (as band-merged with the MIPS sources) were included in our final catalog. Our MIPS map covers more than twice the area of the IRAC map. In the $A_V > 5$ region with full MIPS and IRAC coverage, there are 2608 MIPS-24 point sources, 88% of which have a counterpart in at least one IRAC band; 31 $70 \mu\text{m}$ point sources, all but one of which have an IRAC counterpart; and six

$160 \mu\text{m}$ point sources, all with IRAC counterparts. The wide variation in resolution between IRAC ($\sim 2''$ at $3.6 \mu\text{m}$) and MIPS ($6''$ at $24 \mu\text{m}$, $18''$ at $70 \mu\text{m}$, and $36''$ at $160 \mu\text{m}$) leads to potential confusion of several IRAC objects within a single MIPS beam. Details of the band-merger are given in the c2d delivery documents.¹⁶ For the total MIPS map region, the above source

¹⁶ See the information from R. Cutri et al. at <http://irsa.ipac.caltech.edu/applications/Gator/>.

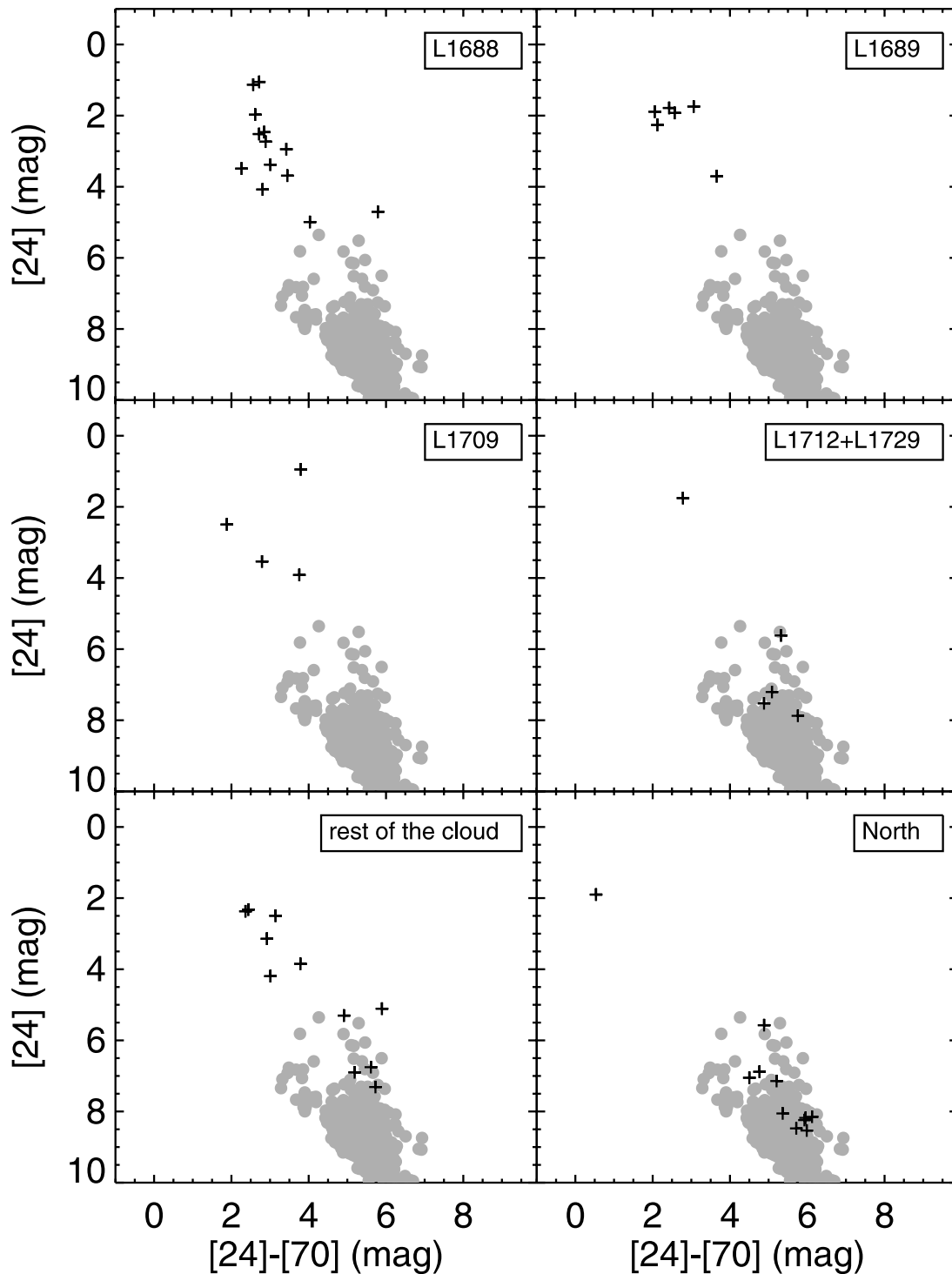


FIG. 12.—A $[24]$ vs. $[24] - [70]$ color-magnitude diagram for Ophiuchus (crosses), with data for SWIRE (gray circles) included for comparison. As in Fig. 11, the two clumps correspond to extragalactic ($[24] \sim 7-10$) and YSO candidates ($[24] \lesssim 5$), and there are few extragalactic objects in the clusters because the high nebulosity limits the depth of the survey. Very red sources ($[24] - [70] > 6$) are the most embedded objects.

counts approximately double (see Table 3). The c2d source catalog containing all Oph *Spitzer* sources can be found online.¹⁷

Near-IR JHK_s data from 2MASS¹⁸ were also used in constructing our catalog. The 2MASS data obviously covered the entire region of the MIPS map, but to a shallower depth than the IRAC observations, so only 45% of the MIPS-24 sources have 2MASS

K counterparts, similar to the fraction seen in the Cham II catalog (Young et al. 2005).

3.4. Color-Magnitude and Color-Color Plots

For relatively bright sources, K_s versus $K_s - [24]$ can be used to separate potential YSOs from stars and most extragalactic sources. This metric has been used in a variety of *Spitzer* papers to identify stars with infrared excess (Gorlova et al. 2004; Padgett et al. 2004, 2006; Rebull et al. 2007; etc.). Stellar photospheres

¹⁷ See http://data.spitzer.caltech.edu/popular/c2d/20051220_enhanced_v1/Oph/.

¹⁸ See the c2d delivery documents by N. J. Evans II et al. online.

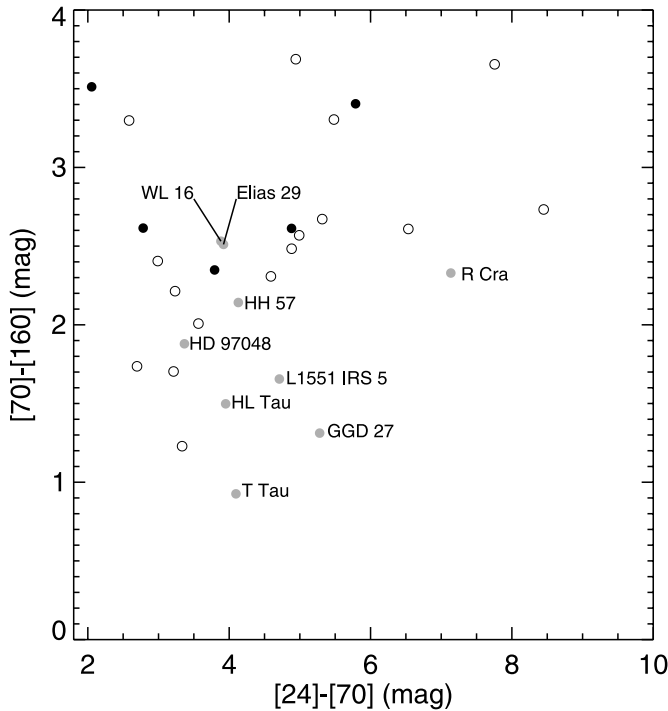


FIG. 13.—Color-color diagram for sources detected but not saturated in all three MIPS bands. Objects shown are from the current Ophiuchus survey (*black circles*), the c2d MIPS Perseus survey (*open circles*; Rebull et al. 2007), and bright objects studied by ISO (*gray circles*; A. Noriega-Crespo 2005, private communication).

without excess are clustered around $K_s - [24] = 0$. If $K_s - [24] > 2$, then a strong excess is present at $24 \mu\text{m}$. Most extragalactic sources are fainter than $K_s = 14$.

Figure 10 presents K_s versus $K_s - [24]$ color-magnitude diagrams for L1688, L1689, L1709, the L1712+L1729 filament, Oph North, the rest of the cloud, and the extragalactic SWIRE ELAIS N1 field. Note that sources that appear to be pure photospheres appear in gray. SWIRE $24 \mu\text{m}$ sources tend to congregate in two groups, namely, main-sequence photospheres with colors of zero and galaxies with faint K magnitudes and colors of 6–8. Thus, $K_s - [24]$ plots help to define regions of color-magnitude space which are likely to be affected by extragalactic sources seen through the cloud. In Figure 10 dashed lines (*right to left*) denote the divisions between Class I, flat, Class II, and Class III objects following Rebull et al. (2007); in brief, we can use the observed $K_s - [24]$ colors and assign an α -index following Greene et al. (1994). Note that the formal Greene et al. classification puts no lower limit on the colors of Class III objects (thereby including those with SEDs resembling bare stellar photospheres, and allowing for other criteria to define youth). In our case, since we know little about many of these objects, in an attempt to limit contamination from foreground/background stars (and background galaxies) we have imposed the additional constraints that $K_s - [24] > 2$ and $K_s < 14$. Note that there are certainly true young cluster members that do not meet these (conservative) criteria, but identifying them (and separating them from the galaxies) is beyond the scope of this paper. L. E. Allen et al. (2008, in preparation) will discuss the combined IRAC and MIPS SEDs and present a list of cloud members.

Interstellar and circumstellar reddening could plausibly affect the location of sources on the K_s versus $K_s - [24]$ diagram. To test this possibility, we used the extinction law derived from *Spitzer* lines of sight through Ophiuchus (Flaherty et al. 2007) to produce

TABLE 5
IRAS RESULTS IN THE MIPS OPHIUCHUS MAP (MAIN+NORTH)

Item	At $24 \mu\text{m}$	At $70 \mu\text{m}$
12 μm PSC real (<i>IRAS</i> qual=3) detections	100	91
Cleanly retrieved.....	74 (74%)	25 (27%)
Completely missing	2 (2%)	41 (45%)
Nebulosity; no point source	15 (15%)	24 (26%)
Resolved as multiple.....	9 (9%)	1 (1%)
25 μm PSC real (<i>IRAS</i> qual=3) detections	63	63
Cleanly retrieved.....	41 (65%)	26 (41%)
Completely missing	1 (1%)	12 (19%)
Nebulosity; no point source	15 (23%)	24 (38%)
Resolved as multiple.....	6 (9%)	1 (1%)
60 μm PSC real (<i>IRAS</i> qual=3) detections	93	84
Cleanly retrieved.....	23 (24%)	20 (23%)
Completely missing	41 (44%)	16 (19%)
Nebulosity; no point source	27 (29%)	48 (57%)
Resolved as multiple.....	2 (2%)	0 (0%)
100 μm PSC real (<i>IRAS</i> qual=3) detections	55	50
Cleanly retrieved.....	4 (7%)	3 (6%)
Completely missing	29 (52%)	12 (24%)
Nebulosity; no point source	22 (40%)	35 (70%)
Resolved as multiple.....	0 (0%)	0 (0%)
12 μm FSC real (<i>IRAS</i> qual=3) detections	50	35
Cleanly retrieved.....	42 (84%)	3 (8%)
Completely missing	0 (0%)	24 (68%)
Nebulosity; no point source	7 (14%)	8 (22%)
Resolved as multiple.....	1 (2%)	0 (0%)
25 μm FSC real (<i>IRAS</i> qual=3) detections	14	11
Cleanly retrieved.....	13 (92%)	5 (45%)
Completely missing	0 (0%)	6 (54%)
Nebulosity; no point source	0 (0%)	0 (0%)
Resolved as multiple.....	1 (7%)	0 (0%)
60 μm FSC real (<i>IRAS</i> qual=3) detections	0	0
100 μm FSC real (<i>IRAS</i> qual=3) detections	0	0

a reddening vector in Figure 10. The vector depicts the effects of 5 mag of K -band extinction, which moves a star toward the lower right. Note that unlike the SWIRE sources, which show a well-defined main sequence at $K_s - [24] = 0$, almost all of the Ophiuchus sources show a displacement toward redder colors. Given the high extinctions throughout this region, many of the apparently small excess sources in L1712+L1729 and the “rest of cloud” areas may simply be reddened photospheres. In addition, Class II sources could plausibly be reddened into the Class I zone. Thus, in regions of very high extinction, classification of YSOs based on even long-wavelength MIPS data is complicated by reddening effects.

Table 4 provides statistics on the number of sources classified by $K_s - [24]$ as YSOs in the various subregions of the Ophiuchus map. The Oph North and off-cloud fields resemble the SWIRE regions, with most sources clumped in the photospheric and faint galaxy region of the diagram. The L1712+L1729 filament shows a few more bright photospheres with small excesses than Oph North plus three potential Class II sources. L1709 and L1689 have a higher frequency of Class II sources than the filament, as well as a few flat and Class I SEDs. Although the number of sources in L1709 is small, this region actually has the highest fraction of Class I + flat spectrum sources in the region. These dense clouds also show fewer sources in the extragalactic clump at faint K_s and large $K_s - [24]$ due to their small area and bright $24 \mu\text{m}$ diffuse emission. Finally, the L1688 region has a high fraction of flat and Class I SED sources (27%), and it dominates the region for sheer number of Class II sources (100).

TABLE 6
IRAS PSC DETECTIONS NOT RECOVERED BY *Spitzer* MIPS

PSC Name	At 24 μm	At 70 μm
16169–2443	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16172–2503	Missing (upper limit at 12, 15)	Confused by nebulosity
16173–2422	Confused by nebulosity	Missing (detected at 60, upper limit at 100)
16184–2452	Confused by nebulosity	Confused by nebulosity
16184–2452	Confused by nebulosity	Confused by nebulosity
16187–2339	Confused by nebulosity	Missing (upper limit at 60, 100)
16193–2335	Confused by nebulosity	Confused by nebulosity
16193–2401	Confused by nebulosity	Confused by nebulosity
16193–2450	Confused by nebulosity	Confused by nebulosity
16194–2410	Missing (upper limit at 12, 25)	Confused by nebulosity
16194–2500	Off edge	Confused by nebulosity
16196–2443	Missing (upper limit at 12, 25)	Confused by nebulosity
16200–2251	Confused by nebulosity	Confused by nebulosity
16201–2410	Retrieved (upper limit at 12)	Confused by nebulosity
16202–2356	Missing (upper limit at 12, 25)	Confused by nebulosity
16202–2427	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16203–2332	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit 100)
16205–2308	Missing (upper limit at 12, 25)	Confused by nebulosity
16209–2422	Confused by nebulosity	Confused by nebulosity
16212–2316	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit 100)
16214–2302	Missing (upper limit at 12, 25)	Confused by nebulosity
16214–2345	Confused by nebulosity	Confused by nebulosity
16214–2436	Confused by nebulosity	Confused by nebulosity
16217–2244	Missing (upper limit at 12, 25)	Confused by nebulosity
16219–2344	Confused by nebulosity	Confused by nebulosity
16219–2417	Confused by nebulosity	Confused by nebulosity
16221–2428	Confused by nebulosity	Confused by nebulosity
16222–2358	Confused by nebulosity	Confused by nebulosity
16223–2404	Confused by nebulosity	Confused by nebulosity
16225–2417	Confused by nebulosity	Confused by nebulosity
16226–2319	Broken into pieces	Confused by nebulosity
16226–2420	Confused by nebulosity	Confused by nebulosity
16227–2418	Confused by nebulosity	Confused by nebulosity
16228–2432	Confused by nebulosity	Confused by nebulosity
16229–2413	Confused by nebulosity	Confused by nebulosity
16233–2421	Confused by nebulosity	Confused by nebulosity
16234–2436	Confused by nebulosity	Confused by nebulosity
16234–2436	Confused by nebulosity	Confused by nebulosity
16235–2416	Retrieved	Confused by nebulosity
16238–2317	Missing (upper limit at 12, 25)	Confused by nebulosity
16246–2440	Retrieved	Confused by nebulosity
16257–2533	Missing (upper limit at 12, 25)	Confused by nebulosity
16262–2545	Retrieved	Missing (detected at 60, upper limit at 100)
16265–2350	Missing (upper limit at 12, 25)	Confused by nebulosity
16267–2539	Confused by nebulosity	Confused by nebulosity
16268–2533	Missing (upper limit at 12, 25)	Missing (detected 60, upper limit at 100)
16269–2454	Confused by nebulosity	Confused by nebulosity
16272–2541	Confused by nebulosity	Confused by nebulosity
16274–2349	Confused by nebulosity	Confused by nebulosity
16275–2251	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16277–2356	Confused by nebulosity	Missing (detected at 60, upper limit at 100)
16280–2353	Confused by nebulosity	Missing (detected at 60, upper limit at 100)
16281–2514	Missing (upper limit at 12, 25)	Confused by nebulosity
16285–2519	Missing (upper limit at 12, 25)	Confused by nebulosity
16289–2527	Missing (upper limit at 12, 25)	Confused by nebulosity
16301–2525	Missing (upper limit at 12, 25)	Confused by nebulosity
16303–2428	Confused by nebulosity	Confused by nebulosity
16304–2504	Missing (upper limit at 12, 25)	Confused by nebulosity
16306–2435	Confused by nebulosity	Confused by nebulosity
16311–2419	Confused by nebulosity	Confused by nebulosity
16313–2439	Confused by nebulosity	Confused by nebulosity
16318–2519	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16322–2421	Confused by nebulosity	Confused by nebulosity
16325–2433	Confused by nebulosity	Confused by nebulosity
16330–2424	Missing (upper limit at 12, 25)	Confused by nebulosity
16330–2431	Missing (upper limit at 12, 25)	Confused by nebulosity

TABLE 6—Continued

PSC Name	At 24 μm	At 70 μm
16330–2431	Missing (upper limit at 12, 25)	Confused by nebulosity
16335–2402	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16335–2419	Missing (upper limit at 12, 25)	Confused by nebulosity
16339–2422	Missing (upper limit at 12, 25)	Confused by nebulosity
16341–2413	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16342–2413	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16375–2439	Missing (upper limit at 12, 25)	Confused by nebulosity
16379–2444	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16383–2407	Missing (upper limit at 12, detected at 25)	Missing (upper limit at 60, 100)
16414–2337	Missing (upper limit at 12, 25)	Missing (detected at 60, upper limit at 100)
16417–2155	Confused by nebulosity	Confused by nebulosity
16418–2155	Confused by nebulosity	Confused by nebulosity
16421–2146	Confused by nebulosity	Confused by nebulosity
16426–2129	Confused by nebulosity	Confused by nebulosity
16436–2121	Confused by nebulosity	Confused by nebulosity

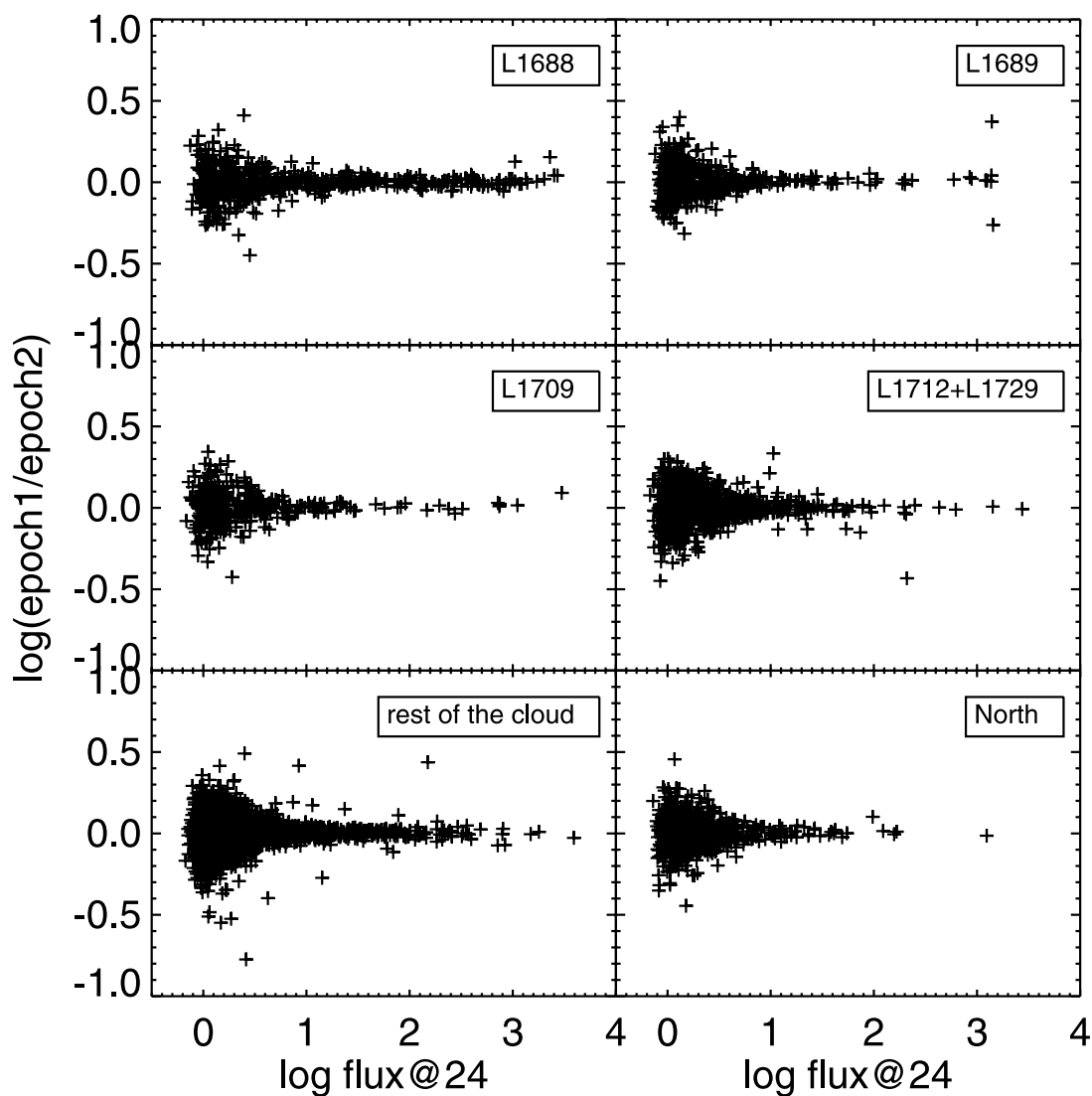


FIG. 14.—Variability of Ophiuchus 24 μm sources between two observations separated by 3–8 hr. All cases of potential variability are well explained by instrumental effects.

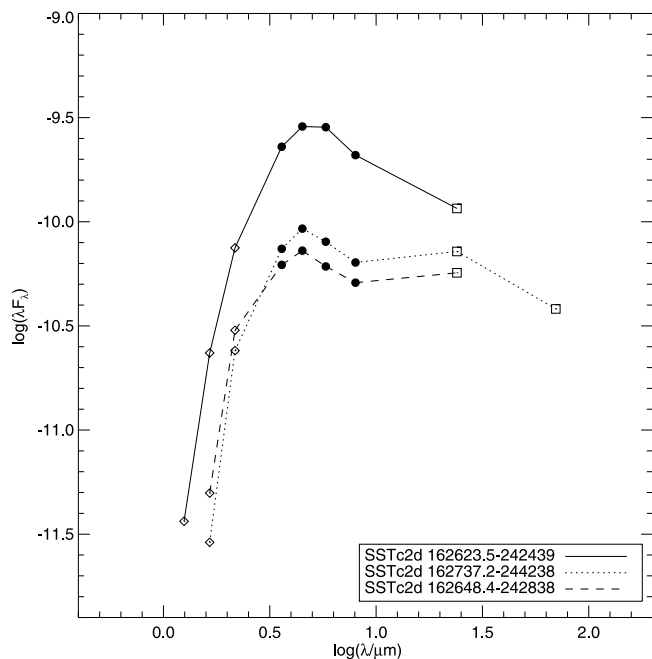


FIG. 15.—Infrared SEDs of several MIPS sources with $K_s - [24] > 8$.

Figure 11 shows the K_s versus $K_s - [70]$ plots for Ophiuchus. In this plot $K_s - [70] = 0$ indicates a bare stellar photosphere. However, few if any bare stellar photospheres in this field are bright enough for detection in our very shallow MIPS $70 \mu\text{m}$ survey. Extragalactic SWIRE sources from the ELAIS N1 field are shown for comparison (*gray circles*). This plot is most useful

for extremely bright YSOs which are saturated at $24 \mu\text{m}$ but not $70 \mu\text{m}$. Again, L1688 dominates the number of potential YSOs and has the most objects with large excesses. These are likely to be the youngest in the sample. Other regions have only a handful of $70 \mu\text{m}$ YSOs, and the bulk are in the Class II regime with the exception of a few sources in L1709 and L1689. There is one object in L1689 with stellar or near-stellar colors out to $70 \mu\text{m}$; it is SSTc2d 163528.6–245648, which we tentatively associate with Elias 2-83. This source may plausibly be a reddened star given the dark cloud visible in this area in the MIPS $24 \mu\text{m}$ image. All the $70 \mu\text{m}$ sources in Oph North are consistent with extragalactic background sources, except for SSTc2d 164417.8–220648, which can be identified with IRAS 16413–2201, a carbon star also known as NC 80 (Guglielmo et al. 1993).

Figure 12 presents a plot of $[24]$ versus $[24] - [70]$ for the Ophiuchus clouds. In the unlikely case of a bare stellar photosphere, $[24] - [70] = 0$. Again, SWIRE sources are depicted as gray circles. Sources with large infrared excesses have large $[24] - [70]$ colors. This color-magnitude diagram shows the separation of YSOs from extragalactic sources and emphasizes that extragalactic $70 \mu\text{m}$ sources are only seen in the lower background regions. Although the number of sources seen in these plots is considerably smaller than in the $24 \mu\text{m}$ -only plots (due to the lower sensitivity of the MIPS $70 \mu\text{m}$ channel), this color-magnitude diagram suggests confirmation of the picture of the L1688 cluster having a lower fraction of very young sources than L1709. However, it is true that a fair number of the most luminous MIPS sources in L1688 are saturated at $24 \mu\text{m}$ ($\geq J_y$), thus eliminating them from these plots. In addition, the small number of sources in L1709 makes the comparison of questionable significance. Interestingly, some low-luminosity sources with large fractional excesses appear to be found outside of the described

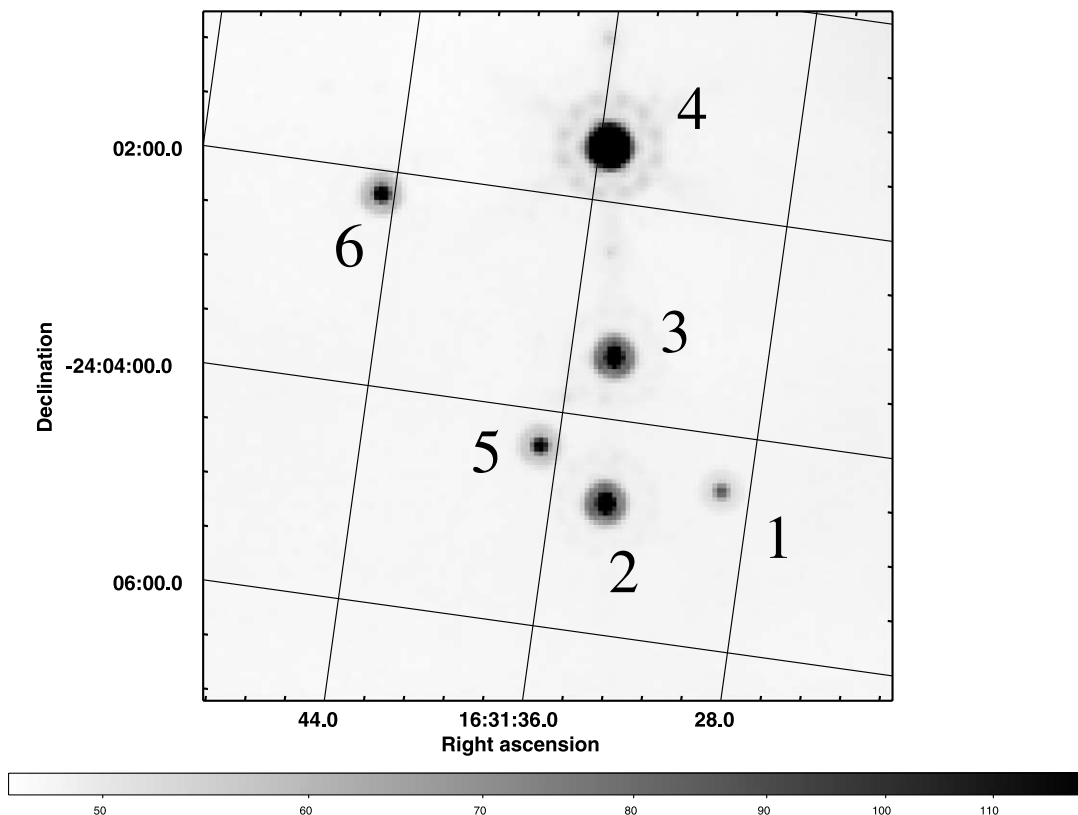


FIG. 16.—A $24 \mu\text{m}$ image and finder chart for the L1709 aggregate.

subclusters (in the “rest of the cloud” areas). Note that the brightest objects in these plots (those with $[24] \leq 5$ and $[24] - [70] \leq 6$) appear for reference in Table 7 below.

Unfortunately, due to very high backgrounds and saturation issues, as well as the difference in sensitivity between the various MIPS arrays, very few point sources (five) have measurable detections at all three MIPS bands. Figure 13 presents a plot of $[70] - [160]$ versus $[24] - [70]$. Included for comparison in this plot are colors calculated for MIPS bands from *ISO* SWS data for a variety of very famous well-studied embedded objects from A. Noriega-Crespo (2005, private communication); to facilitate comparison, these are the same objects used in a similar plot in Rebull et al. (2007). The four Ophiuchus sources all lie between $[70] - [160]$ of 2.5 and 3.5, while spanning a range of 4 mag in $[24] - [70]$. This may be an extinction effect, since three of the five objects are located in regions of Oph with apparently high extinction even at 24 μm .

3.5. Comparison to *IRAS*

As in Rebull et al. (2007) for Perseus, the MIPS data for Oph offer an opportunity to assess how successful the *IRAS* survey was in identifying point sources in a complex region. We now explicitly compare the *IRAS* PSC (Beichman et al. 1988) and *IRAS* Faint Source Catalog (FSC; Moshir et al. 1992) results in this region to the *Spitzer* c2d images and catalogs described in this paper. As for Perseus, extended emission is present in all three MIPS bands, posing a significant source of confusion for the large-aperture *IRAS* measurements. Many of the *IRAS* PSC objects were detected at 60 or 100 μm only, with only upper limits at 12 and 25 μm ; so even without *Spitzer* data, one might suspect that such sources might correspond to texture in the extended emission. As for Perseus (Rebull et al. 2007), the *Spitzer* data comparison clearly shows that significantly fewer spurious sources appear in the FSC than in the PSC.

A summary of the comparison of MIPS sources to *IRAS* sources is found in Table 5. For each *IRAS* PSC and FSC band, the number of *IRAS* detections are listed, followed by the number of these *IRAS* detections recovered by the MIPS 24 and 70 μm arrays, *IRAS* detections not found by MIPS, *IRAS* detections determined to be nebulosity by MIPS, and *IRAS* detections resolved as multiple by MIPS. It is not surprising that few of the long-wavelength *IRAS* sources are recovered by MIPS 24 μm . About 65% of the 25 μm PSC sources are recovered as point sources in MIPS-24; only 23% of the 60 μm PSC sources are recovered at MIPS-70. The most frequent reason for sources not being retrieved is confusion by nebulosity; 23% of the 25 μm PSC sources and 57% of the 60 μm PSC sources are so affected. Surprisingly few objects are resolved into multiple sources by *Spitzer*. Only one 25 μm PSC detection is completely missing from MIPS-24; 16 of the 60 μm PSC detections are completely missing from MIPS-70. A list of *IRAS* PSC sources not cleanly recovered appears in Table 6.

The FSC does a much better job of finding short-wavelength point sources, and is *much* less confused by the nebulosity. However, the FSC does not contain any sources on the west side of the cloud due to high surface brightness in the *IRAS* images. All but one of the 25 μm FSC sources are recovered by *Spitzer* at 24 μm ; the remaining source is resolved as multiple sources when viewed by MIPS. There are no 60 μm FSC high-quality point sources in this region.

While it would be interesting to study long-term far-infrared source variability in Oph, this is very difficult in practice. The vastly larger *IRAS* beam means that point-source flux densities can include large contributions from bright background nebu-

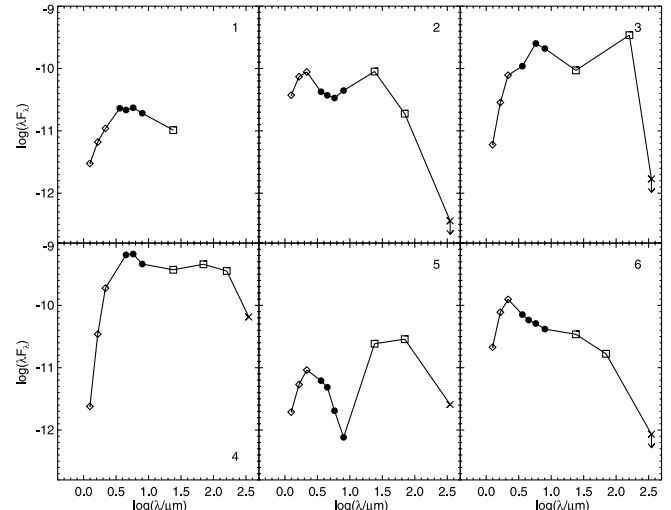


FIG. 17.—SEDs for six stars bright at 24 μm found in the newly identified L1709 aggregate.

losity. The different calibration approaches of the two instruments require that any flux density comparison be made only after large color corrections are applied—corrections that can only be made accurately if the spectral slope is known across the photometric bandpasses. In addition, many sources detected by *IRAS* at 25 μm are near the upper end of the MIPS 24 μm dynamic range, where instrument calibration is less secure. For these reasons, we make no attempt to compare the MIPS and *IRAS* photometry for the sources detected in both surveys.

3.6. Time Variability

Because we have two epochs of observation at 24 μm , we can look for time variability on timescales of a few hours. Figure 14 shows the ratio of flux densities determined in the two epochs as a function of mean 24 μm flux densities. It can be seen that typical variability is $\sim 10\%$, consistent with our reported measurement uncertainty. There are four points whose flux ratio (logarithmic) is greater than 0.45, and two whose flux ratio is less than -0.5 . Out of these six sources, three fall on the edge of the map, and the apparent variability is clearly due entirely to the low redundancy and signal-to-noise ratio at the ends of scan legs. Three additional objects fall in the L1688 region, where bright sources cause the readout-dependent gain variation known as “jailbars.” Two of the L1688 objects are faint and located right on top of some strong jailbars, and one is deep within the bright, extended emission. Therefore, the apparent variability in all cases is entirely due to instrumental effects.

4. DISCUSSION

4.1. Individual Objects

4.1.1. Young Stellar Objects with Extreme $K_s - [24]$ Colors

YSO modeling indicates that some of the earliest evolutionary states of star/disk systems will be detected by *Spitzer* as point sources with extremely red colors (Robitaille et al. 2006). A search through our catalog for sources with $K_s < 12$ and $K_s - [24] > 8$ yielded eight sources. Their SEDs are presented in Figure 15. One of these objects (SSTc2d 162623.5–242439) had been previously identified as CRBR 2322.1–1754 (Comeron et al. 1993). In that paper the source was detected out to the *Q* band, but was not characterized further. In our data this source is probably confused at 24 μm with GY 21, a flat-spectrum YSO

TABLE 7
MIPS PHOTOMETRY FOR POINT-SOURCE OBJECTS DISCUSSED IN THIS PAPER

SSTc2d Name	Region	MIPS-24 ^a (mJy)	MIPS-70 ^a (mJy)	MIPS-160 ^a (mJy)	Other Names ^b	Notes
Red Sources Selected by a Variety of Means						
162145.1–234231	Rest of Cloud	206	730	...	[L89] R7	
162218.5–232148	Rest of Cloud	804	770	...	IRAS 16193–2314=WSB 12	
162309.2–241704	Rest of Cloud	715	1400	...	IRAS 16201–2410	
162502.0–245932	Rest of Cloud	395	630	...	IRAS 16220–2452=WSB 19	
162510.5–231914	Rest of Cloud	840	860	...	IRAS 16221–2312=V2503 Oph=DoAr 16=WSB 20=HBC 257	
162524.3–242756	L1688	127	HD 147889	Discussed in § 4.1.1
162623.5–242439	L1688	927	CRBR 2322.1–1754 (=possibly ISO–Oph 37)	Discussed in § 4.1.1
162646.4–241200	L1688	287	250	...	YLW 37	
162648.4–242838	L1688	455	CRBR 2346.8–2157 (=possibly ISO–Oph 70=WL 2)	Discussed in § 4.1.1
162706.7–243814	L1688	2710	3600	...	BBRCG 22=ISO–Oph 103=WL 17	
162713.8–244331	L1688	318	550	...	ISO–Oph 117=WLY 2-32b	
162717.5–242856	L1688	855	ISO–Oph 124=YLW 12A	Discussed in § 4.1.1
162737.2–244238	L1688	576	890	...	(=possibly ISO–Oph 161 ; 6'' away)	Discussed in § 4.1.1
162738.3–243658	L1688	702	930	...	IRAS 16246–2430=ISO–Oph 163=WLY 1-47=WLY 2-49	
162738.9–244020	L1688	474	1200	...	ISO–Oph 165	
162739.4–243915	L1688	1160	1400	...	ISO–Oph 166=WSB 52	
162739.8–244315	L1688	2520	2900	...	IRAS 16246–2436=ISO–Oph 167=WLY 2-51=WLY 1-45	
162741.6–244644	L1688	72	320	...	ISO–Oph 170	
162816.5–243657	L1688	239	620	...	(Possibly ISO–Oph 196=WSB 60 ; 8'' away)	
162821.6–243623	L1688	94	2100	9910	[SSG2006] MMS 126	
162845.2–242819	L1688	737	1100	...	V853 Oph=DoAr 40=IRAS 16257–2421=ISO–Oph 199=WSB 62	
162857.8–244054	L1688	167	240	...	(Possibly BKLT J162858–244054 ; 7'' away)	
163130.8–242439	L1689	1250	900	4690	IRAS 16284–2418=DoAr 43=WSB 71	
163133.4–242737	L1689	1430	2600	...	IRAS 16285–2421=DoAr 44=SVS 2114=HBC 268=WSB 72	
163133.8–240446	L1709	717	440	...	IRAS 16285–2358	
163135.6–240129	L1709	2980	10647	19000	IRAS 16285–2355=WLY 2-63	
163136.7–240420	L1709	194	670	
163143.7–245524	L1689	1210	1400	...	ISO–Oph 200	
163144.5–240212	L1709	275	390	
163152.4–245536	L1689	887	680	...	ISO–Oph 203	
163154.7–250323	L1689	1380	1400	...	WSB 74 (=possibly IRAS 16289–2457)	
163222.6–242832	L1689	595	IRAS 16293–2422	Discussed in § 4.1.1
163355.6–244204	L1689	235	740	...	RX J1633.9–2442=[NWA2006] SMM 17	
163528.6–245648	L1689	...	740.0	...	CD –24 12715	Debris disk candidate; see § 4.1.1
163945.4–240203	L1712+L1729	1420	2000	4560	IRAS 16367–2356=WSB 82	
164417.8–220648	North	1240	220	...	IRAS 16413–2201	Debris disk candidate; see § 4.1.1
164526.1–250316	Rest of Cloud	150	260	...	IRAS 16424–2457	

TABLE 7—Continued

SSTc2d Name	Region	MIPS-24 ^a (mJy)	MIPS-70 ^a (mJy)	MIPS-160 ^a (mJy)	Other Names ^b	Notes
L1709 Aggregate						
163129.2–240431	L1709	82.2		Source 1 in § 4.1.2
163133.8–240446	L1709	717.0	440.0	...	IRAS 16285–2358	Source 2 in § 4.1.2
163134.2–240325	L1709	747.0	...	18300.0		Source 3 in § 4.1.2
163135.6–240129	L1709	2980.0	10647.0	19000.0	IRAS 16285–2355	Source 4 in § 4.1.2
163136.7–240420	L1709	194.0	670.0	...		Source 5 in § 4.1.2, new edge-on disk candidate
163144.5–240212	L1709	275.0	390.0	...		Source 6 in § 4.1.2
L1689 Aggregate						
163143.7–245559	L1689	19.9		Source 1 in § 4.1.3
163143.7–245524	L1689	1210.0	1400.0	...	ISO–Oph 200	Source 2 in § 4.1.3
163148.9–245429	L1689	20.9	ISO–Oph 201	Source 3 in § 4.1.3
163151.9–245623	L1689	...	1500.0	...	(Possibly ISO–Oph 204=LDN 1689 IRS 5)	Source 4 in § 4.1.3; saturated in nearly all bands
163152.0–245725	L1689	39.8		Source 5 in § 4.1.3; unresolved binary
163152.4–245536	L1689	887.0	680.0	...	ISO–Oph 203	Source 6 in § 4.1.3
163153.4–245505	L1689	104.0	ISO–Oph 205	Source 7 in § 4.1.3
163159.3–245440	L1689	43.8	ISO–Oph 208	Source 8 in § 4.1.3
163200.9–245641	L1689	...	46800.0	...	IRAS 16289–2450=LDN 1689S	Source 9 in § 4.1.3

^a Absolute uncertainties for the 24 μm data are estimated to be 10%–15%; statistical uncertainties are much less than this. Uncertainties for the 70 and 160 μm flux densities are estimated to be 20%.

^b Source associations taken from SIMBAD up to 5'' away from the *Spitzer* position (known to a fraction of an arcsecond); possible associations listed in SIMBAD as being 5''–8'' away are listed as such. In some cases, there are more than a dozen names for these objects, so only the most common are listed. References for these cataloged sources include (L89) Loren (1989), (WSB) Wilking et al. (2005), (BBRCG) Barsony et al. (1989), (WLY) Wilking et al. (1989), (SSG2006) Stanke et al. (2006), (BKLT) Barsony et al. (1997), (NWA2006) Nutter et al. (2006), (DoAr) Dolidze & Arakelyan (1959), and (ISO-Oph) Bontemps et al. (2001).

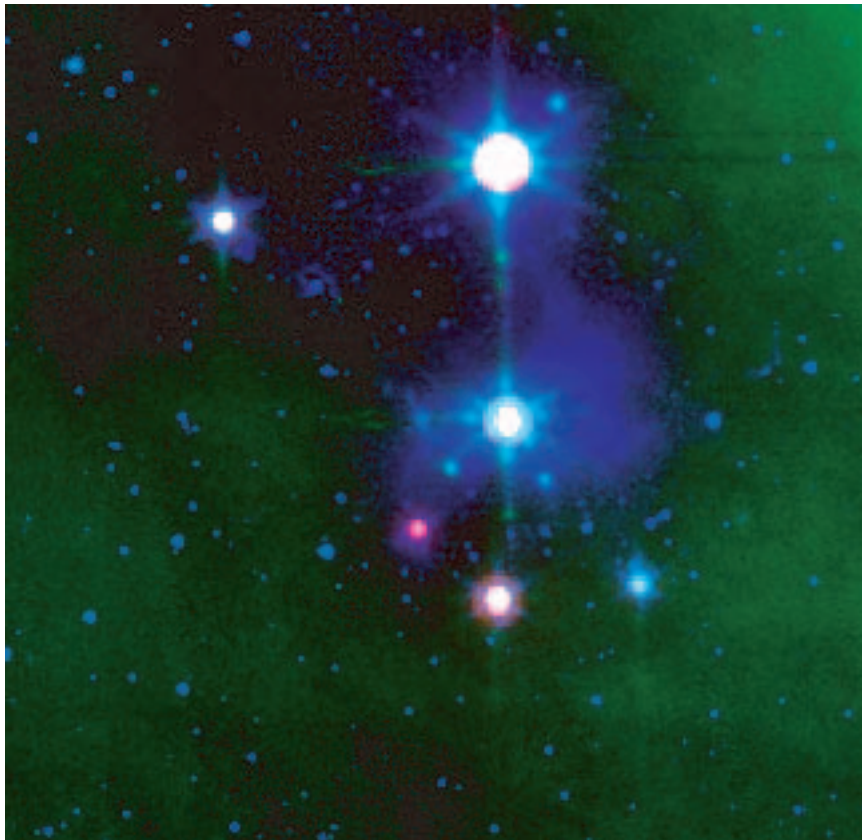


FIG. 18.— Three-color $9' \times 9'$ image of the L1709 aggregate. North is up, and east is left. The IRAC $4.5 \mu\text{m}$ channel is shown in blue, IRAC $8 \mu\text{m}$ in green, and MIPS $24 \mu\text{m}$ in red. The candidate edge-on disk (source 5) is clearly extended in the $4.5 \mu\text{m}$ image.

a few arcseconds away (Greene et al. 1994). Its SED is characteristic of an extinguished stellar photosphere with an excess which appears to be present in the IRAC bands and extends to $24 \mu\text{m}$. No $70 \mu\text{m}$ source is detected. It lies close to another YSO (Oph S2; best seen in our $24 \mu\text{m}$ image), and thus, there is potential SED confusion at longer wavelengths. Source SSTc2d 162648.4–242838 is similar, although fainter. It also appears as a binary source in the IRAC $3.6 \mu\text{m}$ band and can tentatively be identified with WL 2 (Wilking et al. 1989). Finally, SSTc2d 162737.2–244238 (GY 301) is a faint stellar source with an almost flat spectrum out to $70 \mu\text{m}$. All of these sources appear in high-extinction regions, and it appears likely that their extraordinary $K_s - [24]$ colors are primarily due to extinction at the K_s band. Thus, these are likely normal flat-spectrum YSOs with disks seen at very high extinctions due to location within the cloud or envelope.

4.1.2. The L1709 Aggregate and a Possible Edge-on Disk

In L1709 there is a grouping of six bright $24 \mu\text{m}$ point sources within $4.2' \times 3.3'$ (see Fig. 16). SEDs for all six of these objects appear in Figure 17, and MIPS fluxes for these objects appear in Table 7. As a shorthand, numbers between 1 and 6 appear in Figure 17; the corresponding formal c2d names appear in Table 7. IRAS 16285–2358 can be identified with source 2 in Figure 17. This object was selected as a candidate YSO by Beichman et al. (1986). IRAS 16285–2355, source 4 in Fig. 17, is also selected as one of the most extreme $K_s - [24]$ colors in the cloud. This well-known source (e.g., Visser et al. 2002 and references therein) shows a typical Class I or flat-spectrum SED rising through the mid-IR and flattening across the three MIPS bands. It appears to be a version with higher luminosity or more pole-on disk geometry of the same type of sources with extreme $K_s - [24]$ colors

discussed in the previous section. Objects 1, 3, 5, and 6 in Figure 17 do not have any known counterparts in the literature. Source SSTc2d 163136.7–240420, source 5 in Figure 17, stands out for its similarity to edge-on disks models for objects like HH 30 (Burrows et al. 1996). The 2MASS+IRAC part of the SED resembles a normal stellar photosphere which peaks at K , albeit a very faint one. However, its 24 and $70 \mu\text{m}$ fluxes are almost 2 orders of magnitude higher than the $8.0 \mu\text{m}$ value. Edge-on disks such as HH 30 and HK Tau/c (Wood et al. 2002) can appear in scattered light until $10\text{--}20 \mu\text{m}$, suppressing both their short- and long-wavelength flux densities. A $3.6 \mu\text{m}$ image is presented in Figure 18.

Five of the six sources in L1709 were observed in the $350 \mu\text{m}$ continuum at the CSO. Two (L1709-3 and -4) were clearly detected and showed extended structure. For these sources, an elliptical aperture with radius of $1'$ was used to calculate the flux density, excluding some of the low-level extended structure in the images. The L1709-5 (SSTc2d 163136.7) source is in a region of considerable structure; it appears to be sitting on a ridge of emission and shows a cometary structure extending toward the south-southeast. To the northwest ($\alpha = 16^{\text{h}}31^{\text{m}}35.33^{\text{s}}$, $\delta = -24^{\circ}03'46.8''$ [J2000.0]) is another bowed extended ridge of emission, with a flux of ~ 200 mJy. The CSO flux densities and upper limits are incorporated in the SEDs presented in Figure 17. We use the equations of Beckwith et al. (1990) to estimate lower limits on the circumstellar mass from these measurements. Presuming that the $350 \mu\text{m}$ emission is optically thin and $T_d = 100$ K, $\kappa_\nu = 0.071$, leading to lower limits on the circumstellar masses of $1.6 \times 10^{-2} M_\odot$ for SSTc2d 163135.6–240129 (L1709-4) and $6.4 \times 10^{-4} M_\odot$ for SSTc2d 163136.7–240420 (L1709-5).

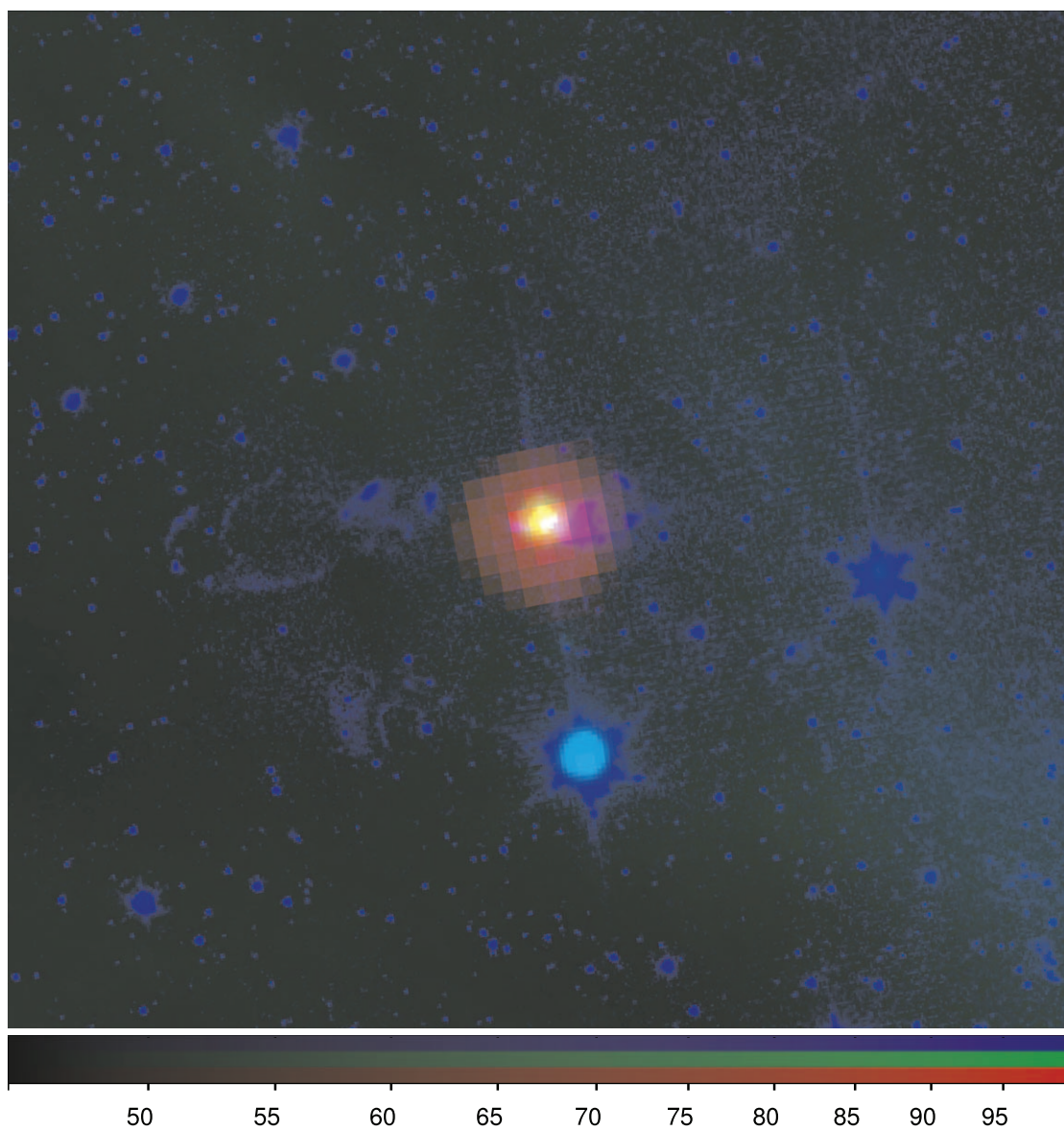


FIG. 19.— Three-color $9' \times 9'$ image of IRAS 16293–2422. The IRAC $8 \mu\text{m}$ channel is shown in blue, MIPS $24 \mu\text{m}$ in green, and MIPS $70 \mu\text{m}$ in red. The source is invisible at $8 \mu\text{m}$, although the outflow emission and reflection nebulae are visible. Note especially the bow-shock features in blue to the left and right of the central source. In this image, north is up, and east is to the left.

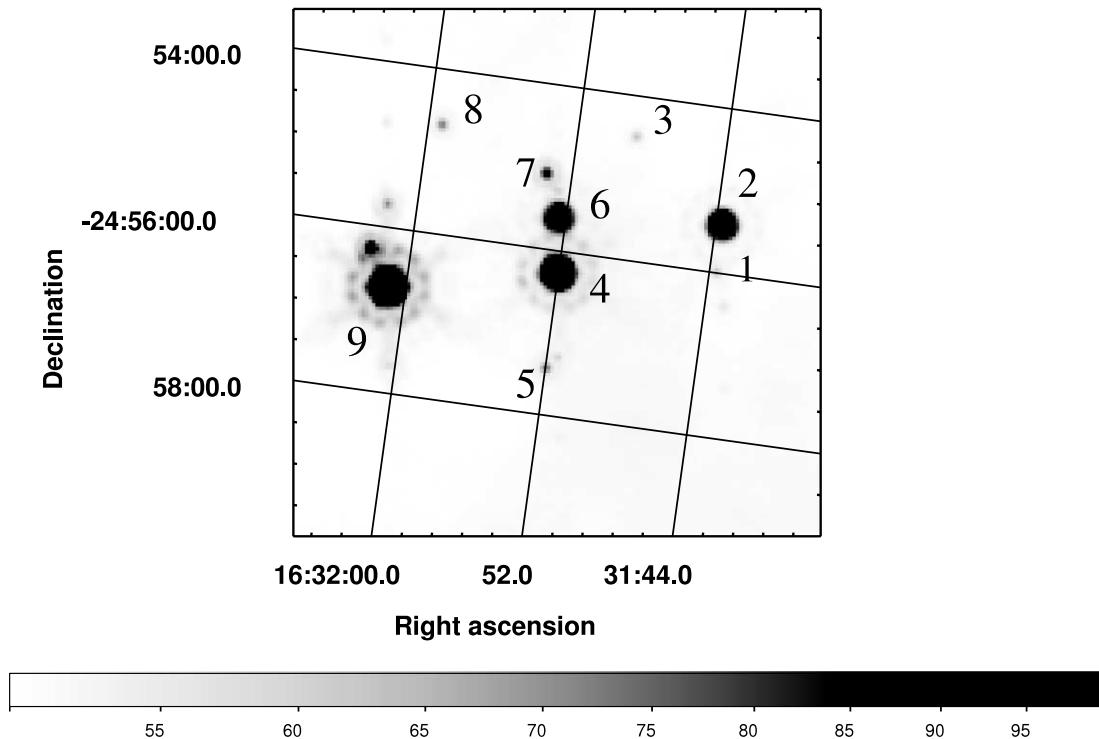


FIG. 20.—A $24\ \mu\text{m}$ image and finder chart for the L1689 aggregate.

4.1.3. L1689

The famous Class 0 source IRAS 16293–2422 appears in the L1689 part of our MIPS maps. Interestingly, the source is entirely invisible in the IRAC $8\ \mu\text{m}$ image, although some of the jet and bow-shock emission from the outflow is plainly visible. The source itself begins to be seen at faint levels at $24\ \mu\text{m}$, is very bright at $70\ \mu\text{m}$, and is saturated at $160\ \mu\text{m}$. A color image is shown in Figure 19. The $8\ \mu\text{m}$ visibility of shocked gas very close to the invisible central source implies that the high-extinction region is very confined in areal coverage, suggesting an edge-on orientation to the disk/envelope system. This source has also been observed by the *Spitzer* Infrared Spectrograph (IRS; Jørgensen et al. 2005). In this work the authors interpret the rapid rise in the SED as a large inner hole in an extended envelope around a protobinary. We prefer to interpret the very large slope in the SED as an extinction effect from a highly inclined disk/envelope system.

There is an aggregate of nine stars bright at $24\ \mu\text{m}$ found clustered within a $5'$ region within L1689. The MIPS- $24\ \mu\text{m}$ image of the aggregate is presented in Figure 20, and the *Spitzer*+2MASS SEDs for these stars appear in Figure 21. Several of these objects were also seen by *ISO* and discussed as part of Bontemps et al. (2001).

Padgett et al. (2004) and Rebull et al. (2007) commented on small clusterings of sources bright at $24\ \mu\text{m}$ found in L1228 and Perseus, respectively. While Oph does not have such aggregates in as high abundance as Perseus, this grouping in L1689 is clearly a similar phenomenon in that there are several sources within $\sim 0.15\ \text{pc}$, all bright at $24\ \mu\text{m}$ and with a diversity of SEDs. As can be seen in Figure 21, some sources resemble reddened photospheres with just a hint of excess at $24\ \mu\text{m}$ (e.g., source 1), some appear to have Class II SEDs (e.g., source 8), and some have substantial excesses like Class I or flat-spectrum SEDs (e.g., source 6). Source 5 has a strangely shaped SED because it is an unresolved binary at $24\ \mu\text{m}$. Source 9 is so bright at $24\ \mu\text{m}$ that it leaves several latent images on the array (e.g., it is near saturation), so the unusual

shape of the IRAC portion of the SED is likely to be a result of saturation in those bands. Object 4 is saturated at all *Spitzer* bands except for $70\ \mu\text{m}$. The proximity of these sources suggests that they are physically associated, despite their disparate SEDs. As in L1228 (Padgett et al. 2004), this aggregate seems to have sources at a variety of stages in their circumstellar evolution, although the sources must be nearly coeval.

4.2. Clustering of YSOs in Ophiuchus

Figure 1 shows the distribution of MIPS $24\ \mu\text{m}$ sources in the c2d maps. The sources appear to be randomly distributed in most of the map. However, as expected, there is a large overdensity of sources in the region around the Oph A core. Figure 22 accentuates this result, showing that relatively bright sources with $24\ \mu\text{m}$ excess are strongly concentrated in the core, as well as a scattering of small groups in L1689, northern L1688, and southern L1709. A modest number of excess sources are more widely distributed over the entirety of the L1719+1729 eastern filament, while Oph North has few bright $24\ \mu\text{m}$ excess sources. The “rest of the cloud,” largely concentrated to the northeast of L1688, shows some intriguing small individual sources but no YSO aggregates.

Table 4 quantifies the degree of clustering seen by categorizing the objects with $24\ \mu\text{m}$ excess. Consideration of all $24\ \mu\text{m}$ sources with 2MASS associations shows that only 297 out of a total of 2590 sources are confined to the L1688 subregion. This would imply a large distributed population of potentially young sources in the cloud. However, when the selection criteria $K_s - [24] > 2$ and $K_s < 14$ are applied, we see that more than half of the YSO candidates are confined to L1688. About 60% of the YSO candidates fall into the L1688, L1689, and L1709 regions. The clustering of sources with Class I and flat SEDs is more pronounced. A full 95% of these sources are confined to the central cluster and the L1709 and L1689 aggregates. In comparison, only 61% of the Class II candidates are located in these subregions. It is worth noting that the natures of the Class I and II candidates

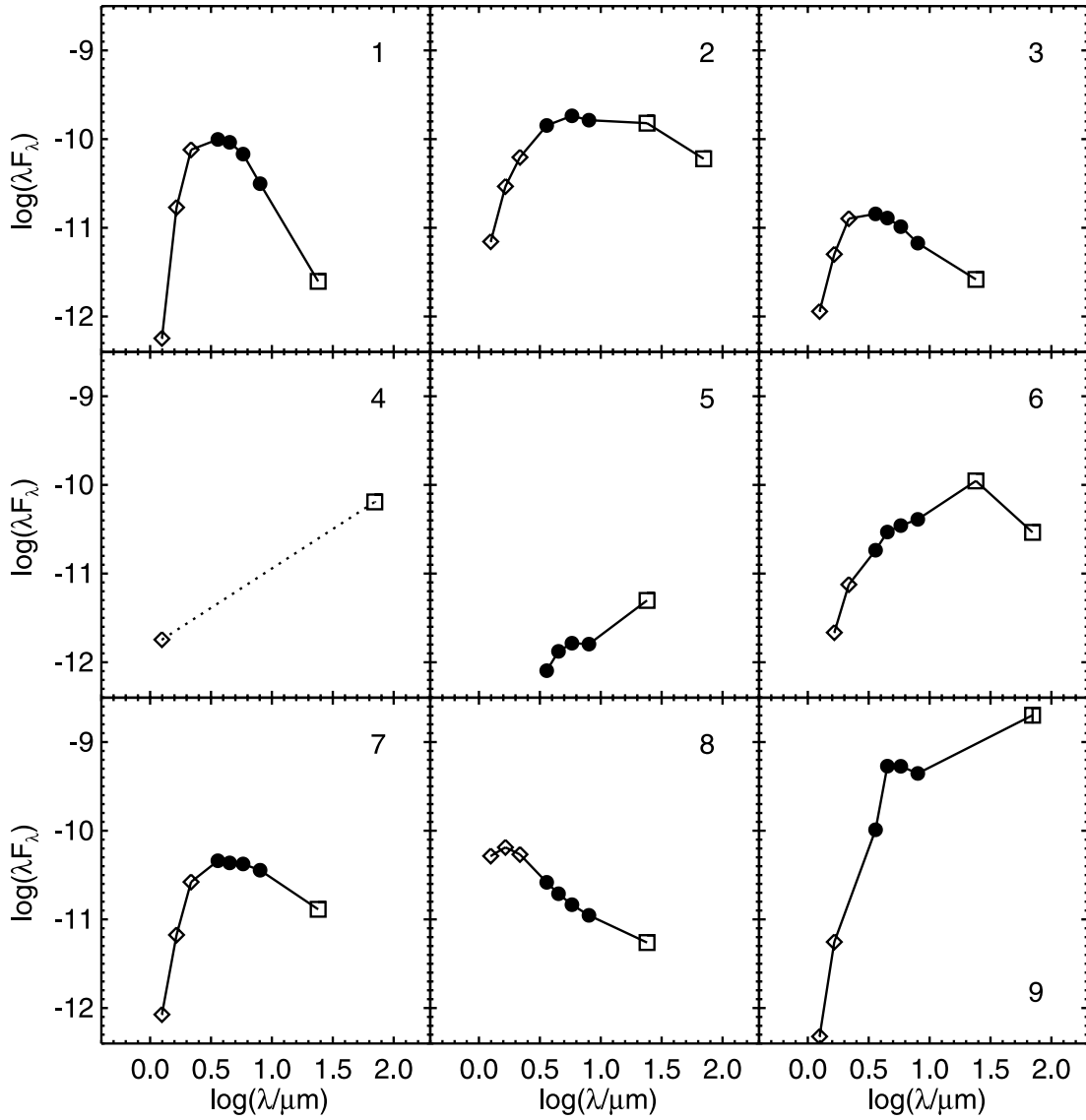


FIG. 21.—SEDs for nine YSO candidates bright at 24 μm found tightly clumped together in L1689. Source 4 is saturated in all *Spitzer* bands except 70 μm .

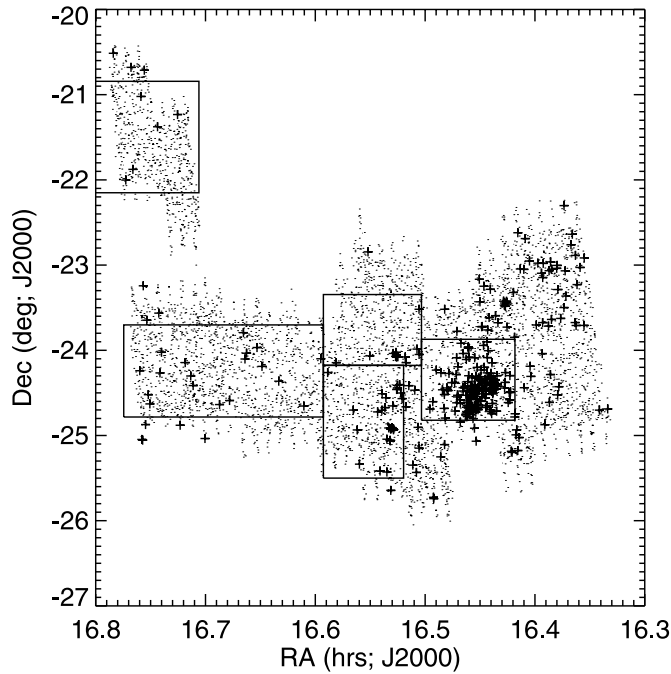


FIG. 22.—Locations of sources where $K_s - [24] > 2$ and $K < 14$.

remain unconfirmed by further study outside of the well-studied cluster and aggregates. Some percentage of these sources may well be extragalactic interlopers. Thus, star formation in the Ophiuchus clouds appears to be largely confined to the cluster and aggregates. Small aggregates of the type seen in Taurus, Lupus, and Perseus are found in the L1709 and L1689 regions, but not in the remainder of the filamentary clouds in Ophiuchus.

5. CONCLUSIONS

We have presented the 14.4 deg^2 c2d *Spitzer* MIPS long-wavelength imaging survey of the Ophiuchus molecular cloud complex. Many point sources and a variety of extended emission structures on scales from circumstellar to intercloud are shown in our mosaicked images. Using point-source-fitting photometry,

we have identified and measured the brightness of 5779 sources at $24 \mu\text{m}$, 81 sources at $70 \mu\text{m}$, and eight sources at $160 \mu\text{m}$. Using criteria derived from a K_s versus $K_s - [24]$ and other color-color and color-magnitude diagrams, we have identified some 323 candidate young stellar objects throughout the clouds. Several sources with extreme colors are discussed, and a candidate edge-on disk is identified. Our survey suggests that although the numbers of forming stars are dominated by the L1688 cluster and small aggregates in L1689 and L1709, there may be a population of Class II sources lurking in the more tenuous parts of the cloud. This distributed population could contain up to 40% of the YSO candidates in Ophiuchus.

Most of the support for this work, part of the *Spitzer Space Telescope* Legacy Science Program, was provided by NASA through contracts 1224608, 1230782, and 1230779, issued by the Jet Propulsion Laboratory, California Institute of Technology, under NASA contract 1407. We thank the Lorentz Center in Leiden for hosting several meetings that contributed to this paper. Support for J. K. J. and P. C. M. was provided in part by a NASA Origins grant, NAG5-13050. Astrochemistry in Leiden is supported by a NWO Spinoza grant and a NOVA grant. K. E. Y. was supported by NASA under grant NGT5-50401 issued through the Office of Space Science. This research has made use of NASA's Astrophysics Data System (ADS) Abstract Service, and of the SIMBAD database, operated at CDS, Strasbourg, France. This research has made use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center, funded by the National Aeronautics and Space Administration and the National Science Foundation. These data were served by the NASA/IPAC Infrared Science Archive, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. The research described in this paper was partially carried out at the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

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